

Pre-supernova neutrino alarm at KamLAND and its extension to an combined system with Super-Kamiokande

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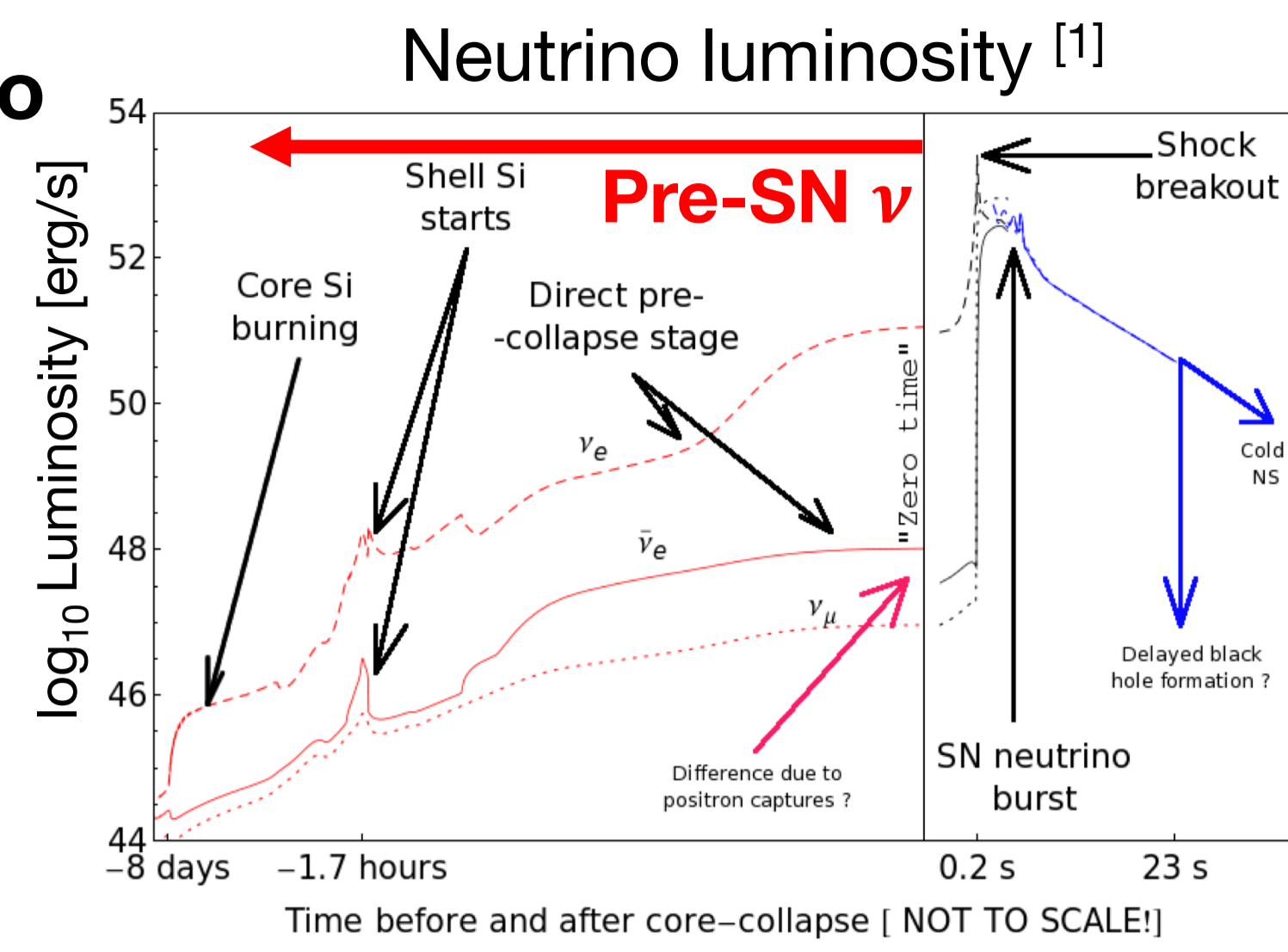
1. Introduction

Pre-supernova (pre-SN) neutrino

These neutrinos are generated by massive star ($M \geq 8M_{\odot}$) at the end of stellar evolution.

Observation of pre-SN neutrinos
→ insight into stellar evolution
→ hints toward understanding neutrino mass ordering

Early warning to SN (Pre-SN alarm)



Kamioka Liquid scintillator Anti-Neutrino Detector (KamLAND)

1000m depth
 $\times 10^{-5}$ muon flux

1000t liquid scintillator (LS) detector

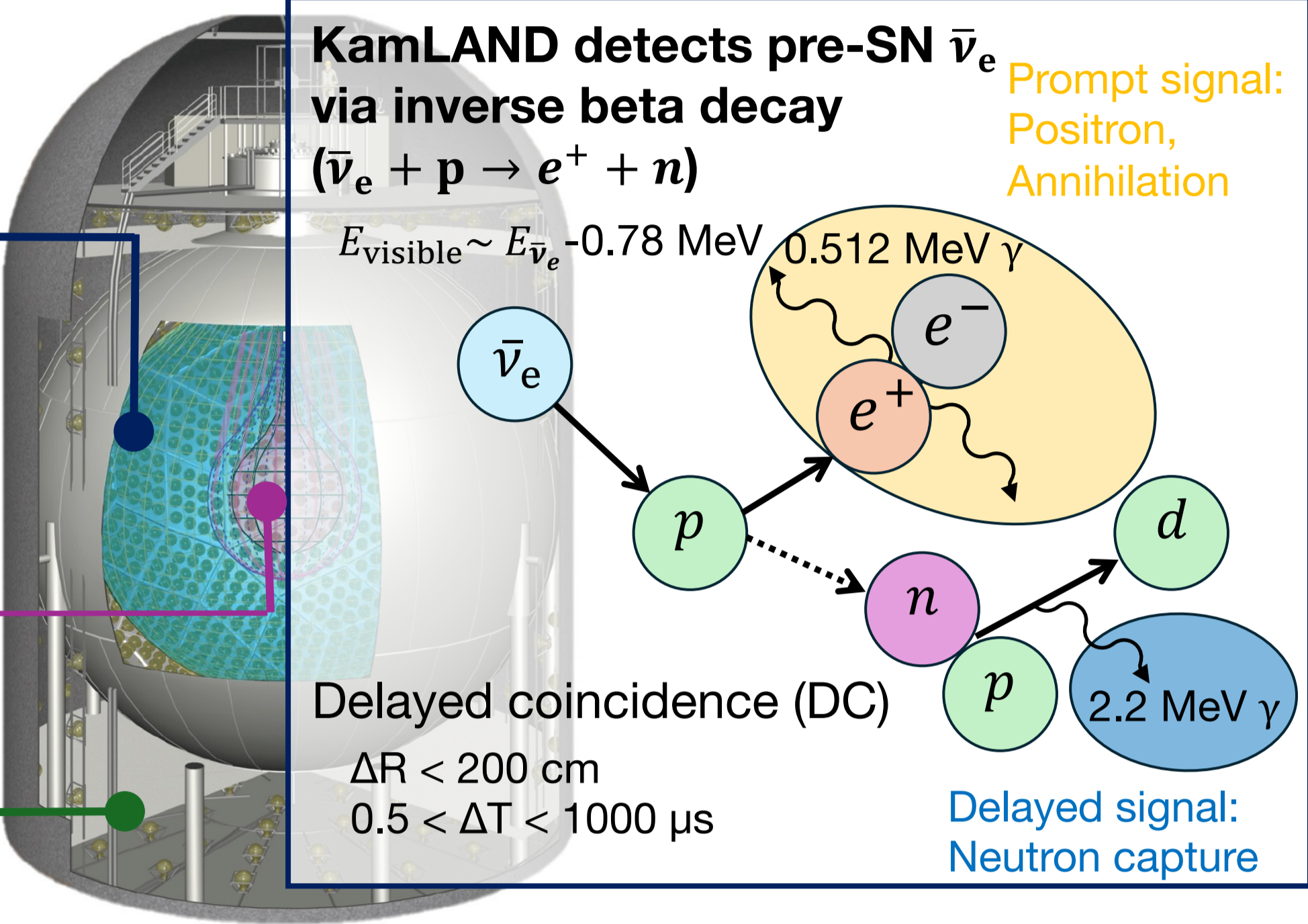
Dodecan (80%), Pseudocumene(20%), PPO (1.36g/l)

with 1325 17-inch PMTs + 554 20-inch PMTs

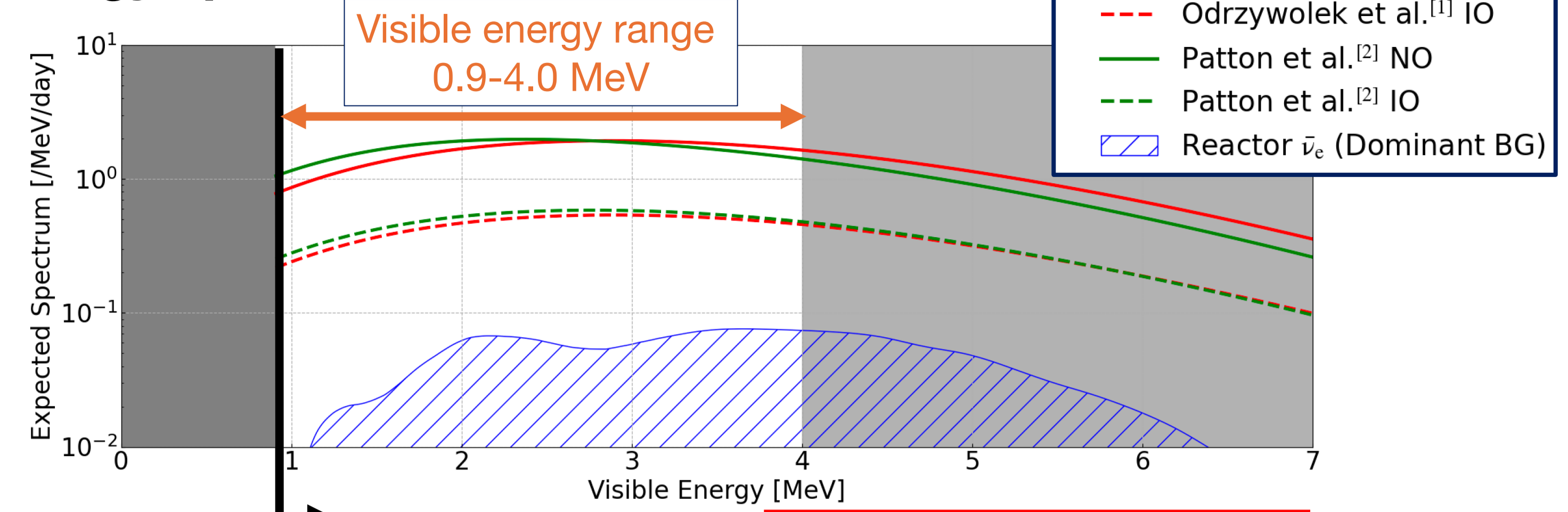
Inner balloon

Xe loaded LS for $0\nu\beta\beta$ search
Volume cut ($r=1.92$ m)

Pure water
For muon veto



Energy spectrum



Low energy threshold $E_{\bar{\nu}_e} > 1.8 \text{ MeV}$

Low BG (0.17 [/day]) with DC scheme

KamLAND has **unique sensitivity to Pre-SN $\bar{\nu}_e$**

Current online Pre-SN alarm system at KamLAND (since 2015) [3]

- The system is based on the significance of statistical excess from BG rate. (**Rate analysis**)
- Time window is 24 hour.
- BG rate is the average one over last 90 days.

Motivation

In order to **improve alarm sensitivity**,

1. **Incorporating time profile**

2. **Combining with Super-Kamiokande**

This poster shows **improved sensitivity**

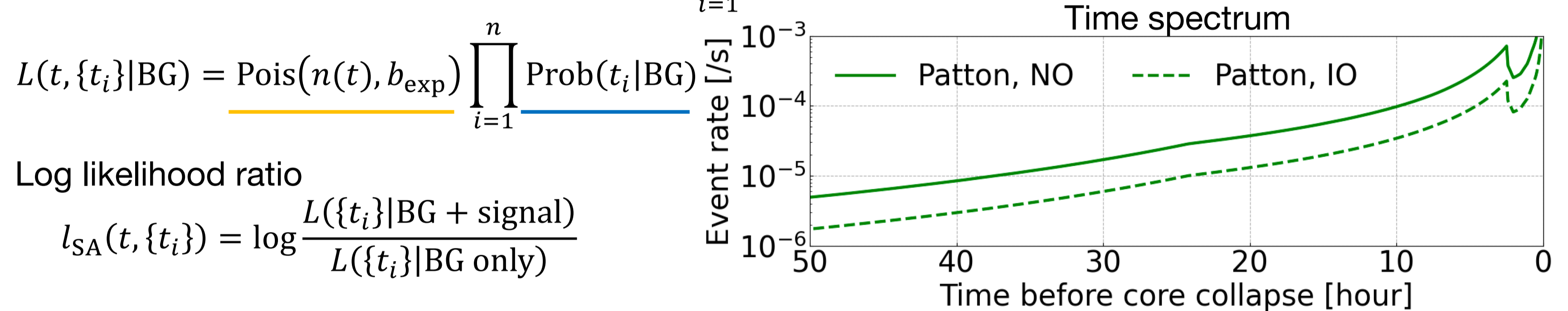
2. Incorporating time profile (Rate+Shape analysis)

Concept of alarm system with time profile of pre-supernova neutrinos [4]

Likelihood function $L(t, \{t_i\} | \text{BG} + \text{Signal}) = \text{Pois}(n(t), s_{\text{exp}} + b_{\text{exp}}) \prod_{i=1}^n \text{Prob}(t_i | \text{BG} + \text{Signal})$

rate term shape term

$\{t_i\}$: event time array ($\subset [t, t + 200 \text{ hour}]$)
($n = \text{number of } \{t_i\} \text{ elements}$)

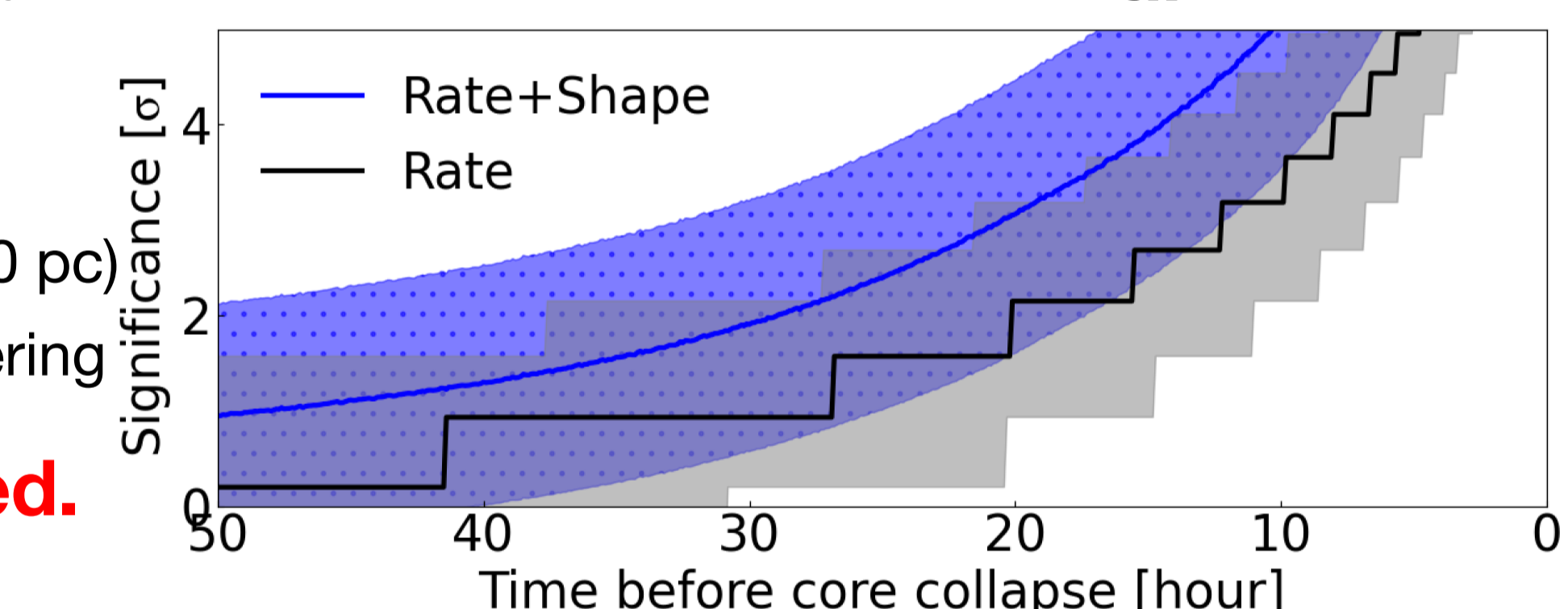


To distinguish between BG and BG+Signal hypotheses, we calculate the threshold of l_{SA} from event time array using toy MC simulation.

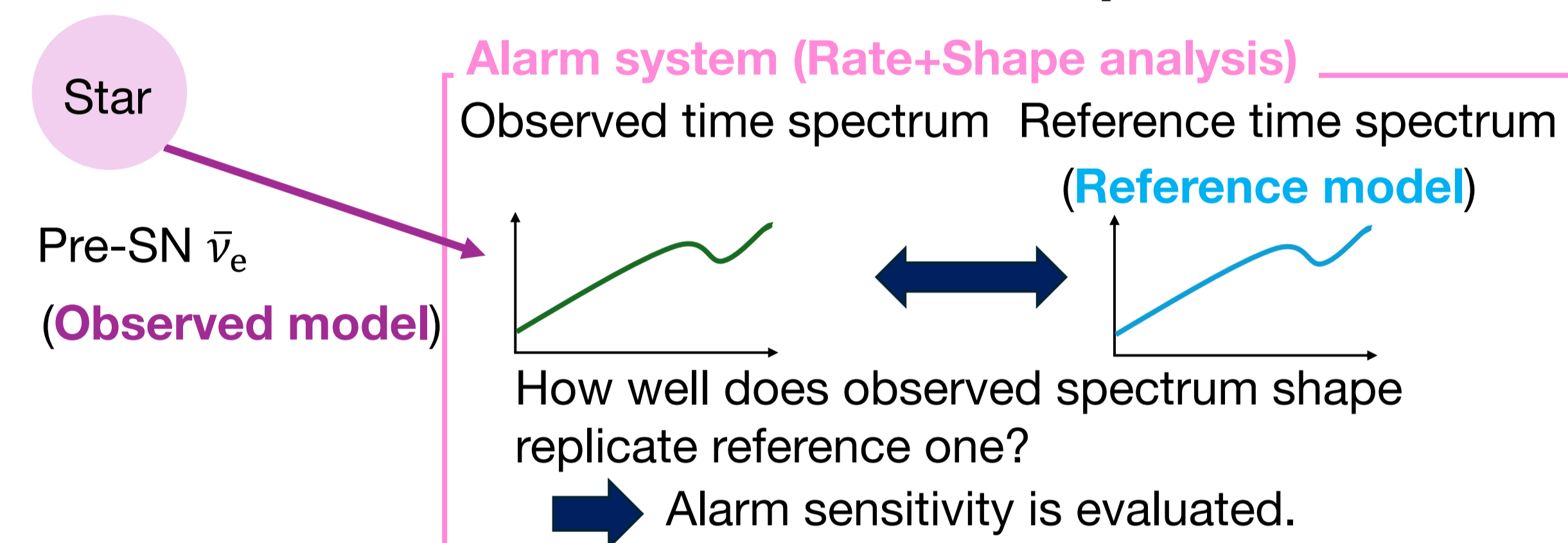
Sensitivity

Target star: Betelgeuse-like star ($15 M_{\odot}$, 150 pc)
Model: Patton model with normal mass ordering

Alarm sensitivity is improved.



Check of the robustness for model dependence



Warning time [hour] (Expected significance $> 3\sigma$)

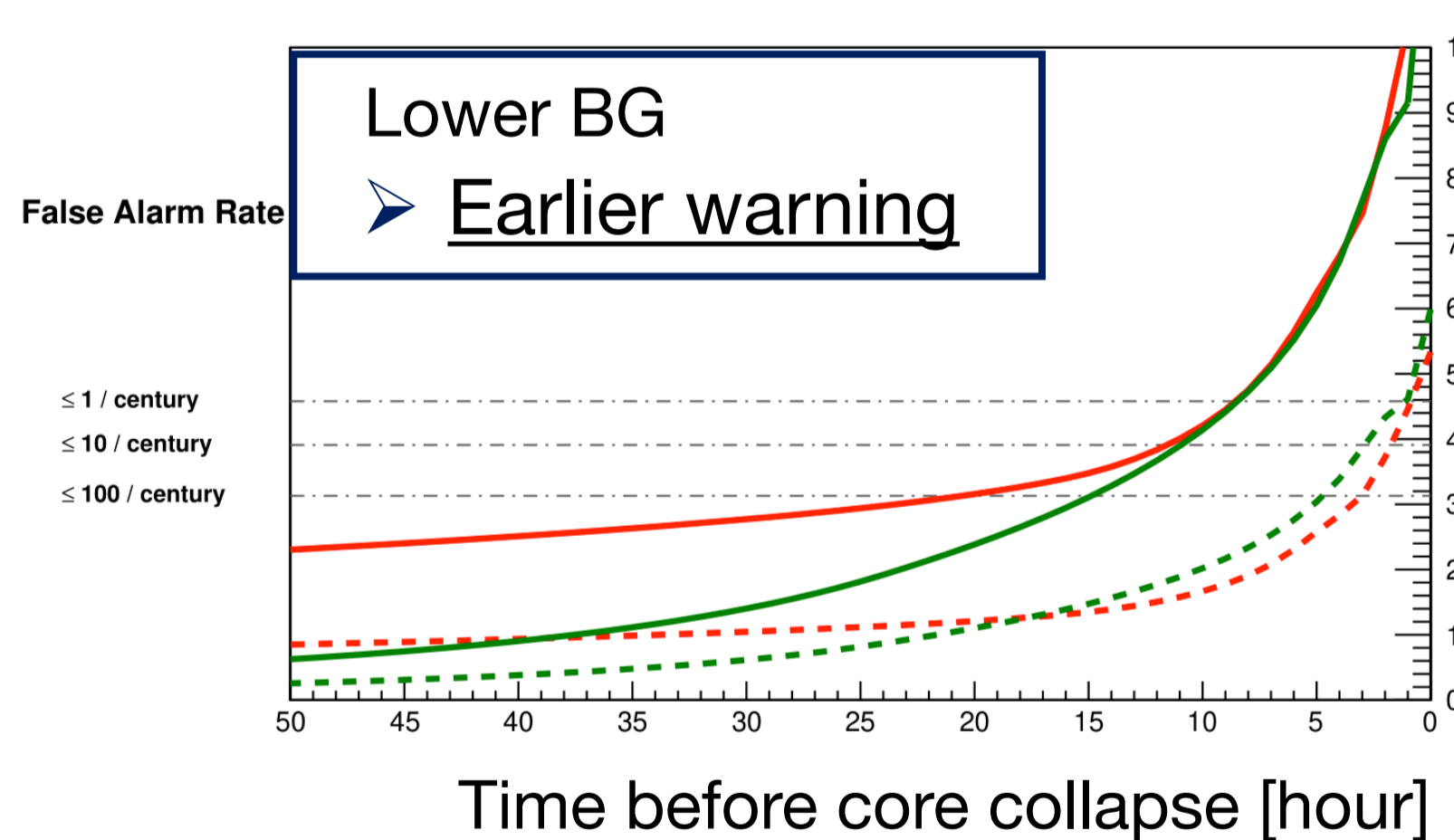
Observed model	Rate+Shape Reference model		Rate
	Patton NO	Patton IO	
Patton NO	20.4	20.2	12.2
Patton IO	7.3	7.3	2.4

Rate+Shape analysis has higher sensitivity than Rate analysis.

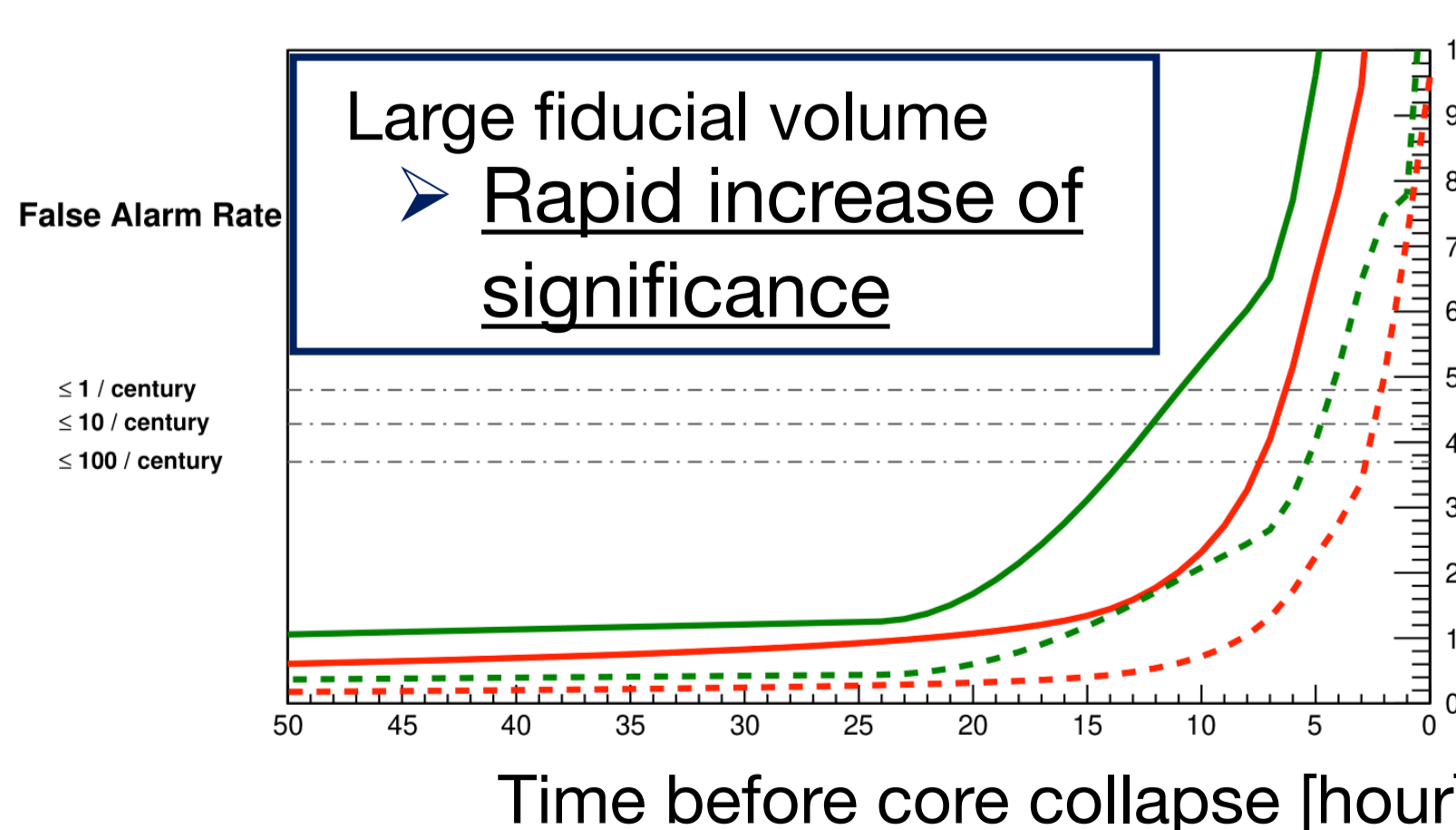
3. Combining with Super-Kamiokande

This analysis uses **only rate**.

KamLAND



SK



Super-Kamiokande (SK)

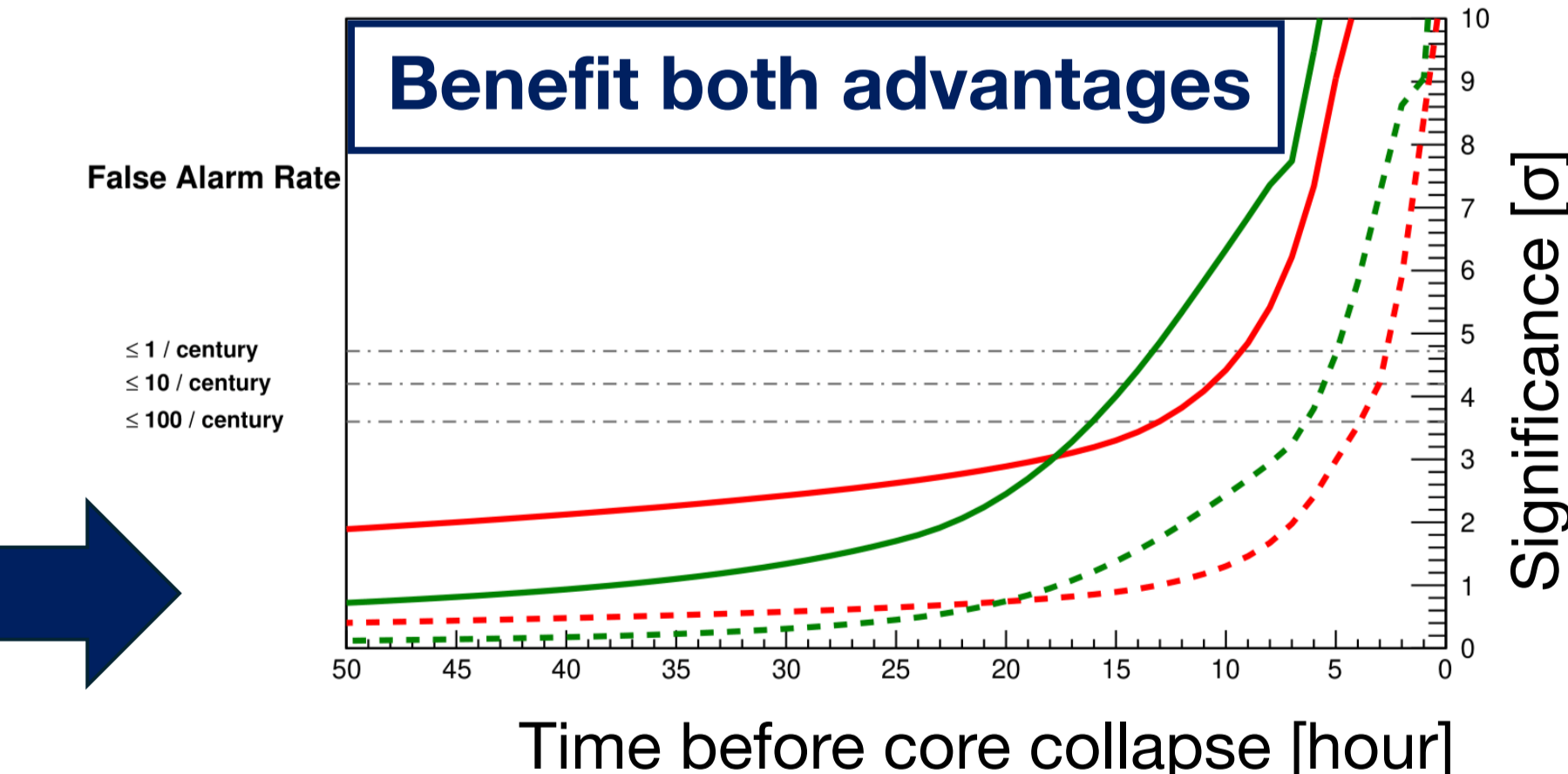
- SK is 22.5kt water Cherenkov detector.
- 0.03% of gadolinium was loaded. (SK-Gd)

➤ **SK can detect pre-SN $\bar{\nu}_e$**

➤ Online pre-SN alarm system has been started since 2021. [5]

Combined of KamLAND & SK

The alert system is triggered based on BG rate of both KamLAND and SK



Warning time [hour]

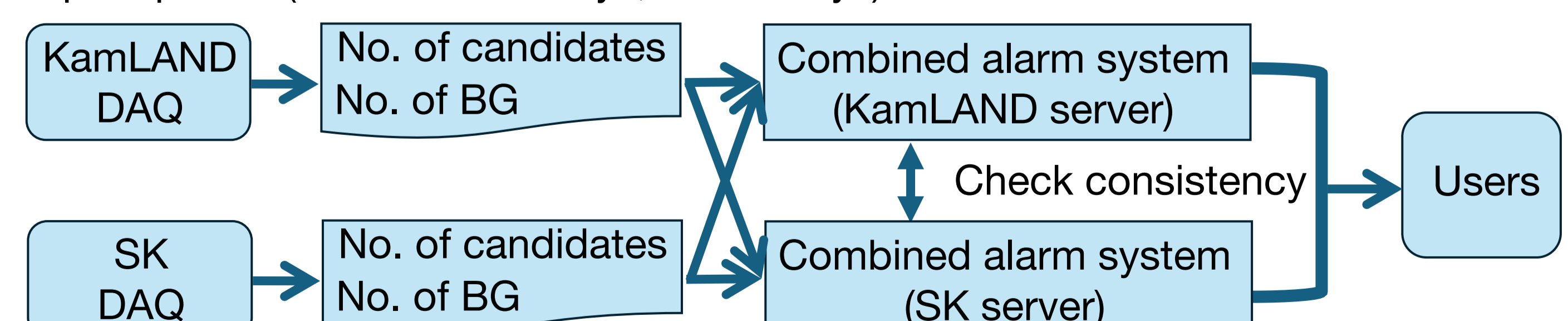
(Expected significance corresponding to FAR $< 1/\text{century}$)

model	Mass ordering	KamLAND	SK	Combined
Odrzywolek	NO	7.4	6.3	8.7
	IO	0.4	2.1	2.6
Patton	NO	7.3	10.9	12.9
	IO	0.7	4.3	4.8

Combined alarm provides earlier warning than individual alarms.

Online combined alarm system (<https://www.lowbg.org/presnalarm/>)

- Running in both KamLAND and SK side (redundancy system)
- BG: Averaged event number over a past period (KamLAND 90days, SK: 30days)
- Total latency time ~ 6 min
- Output every 5 minutes
- Link GCN circular



4. Summary and prospect

- ✓ In order to **improve alarm sensitivity**, this poster presents two approaches.
- ✓ We develop an earlier-issued alarm which use **time spectrum** than the current alarm.
- ✓ In order to **improve alarm sensitivity**, We develop combined alarm system of KamLAND and SK. (using only rate)
- ✓ The combined alarm system is running.

- Study of **estimation the core-collapsed time**
- Development of rate+shape(time)+shape(energy) alarm system
- Submitting of paper about combined pre-SN alarm system

Reference

- [1] Odrzywolek et al., APSS B B41 (2010) 1611
[2] Kelly M. Patton et al (2017) ApJ 851 6
[3] K. Asakura et al. Astrophys.J. 818 (2016) 1, 91
[4] A. Sheshukov et al JCAP12(2021)053
[5] L. N. Machado et al (2022) ApJ 935 40