マルチメッセンジャー天文学のための超新星爆発の長時間計算

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第九回　超新星ニュートリノ

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https://hubblesite.org/resource-gallery

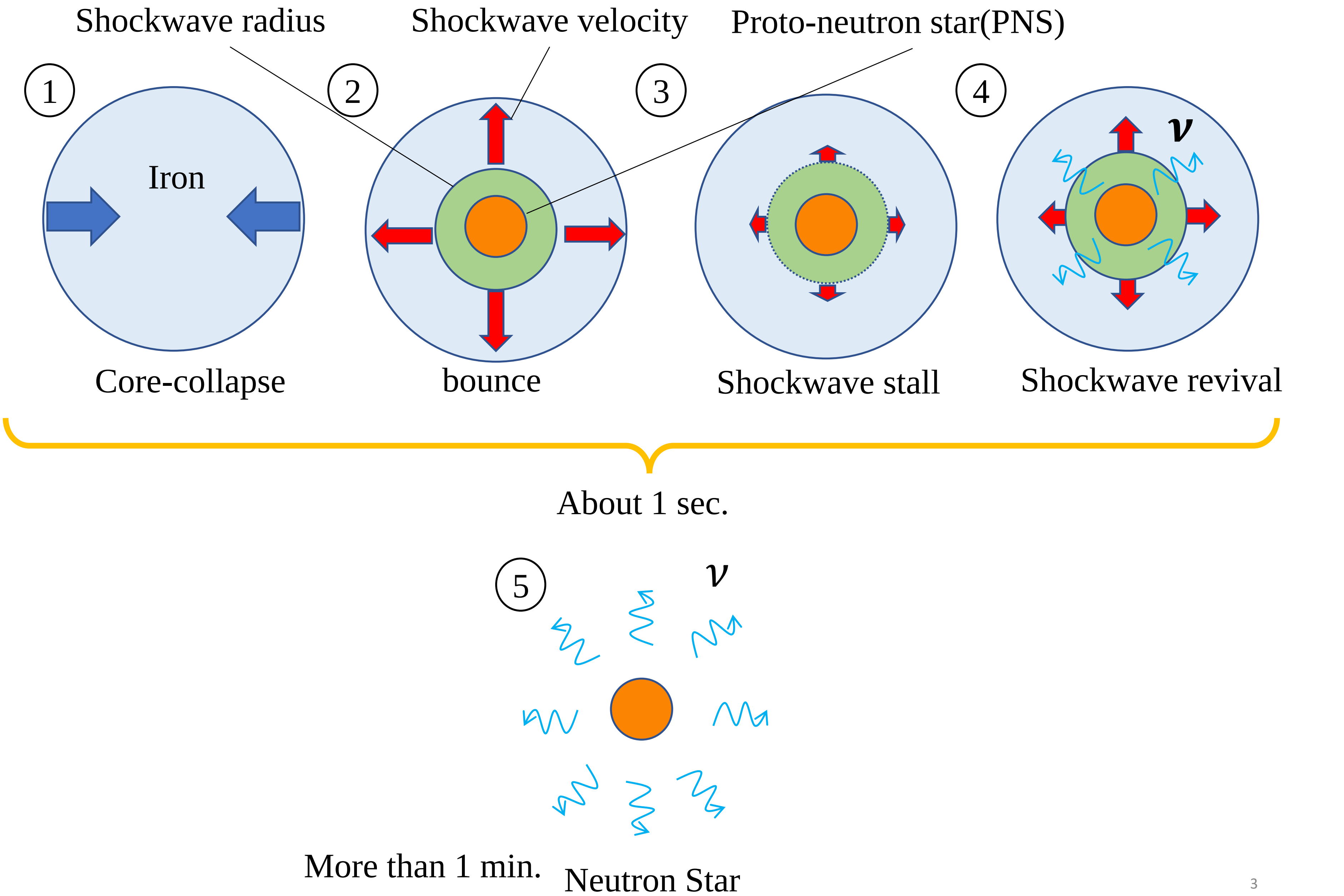
Contents

1. Long-term simulation of supernova neutrino
2. Estimation of neutrino signals on earth
   - arXiv: 2010.16254
3. Estimation of gravitational wave frequencies
   - arXiv: 2302.00292

Keywords

Supernova neutrino, Super-Kamiokande, Neutrino observation, gravitational wave,

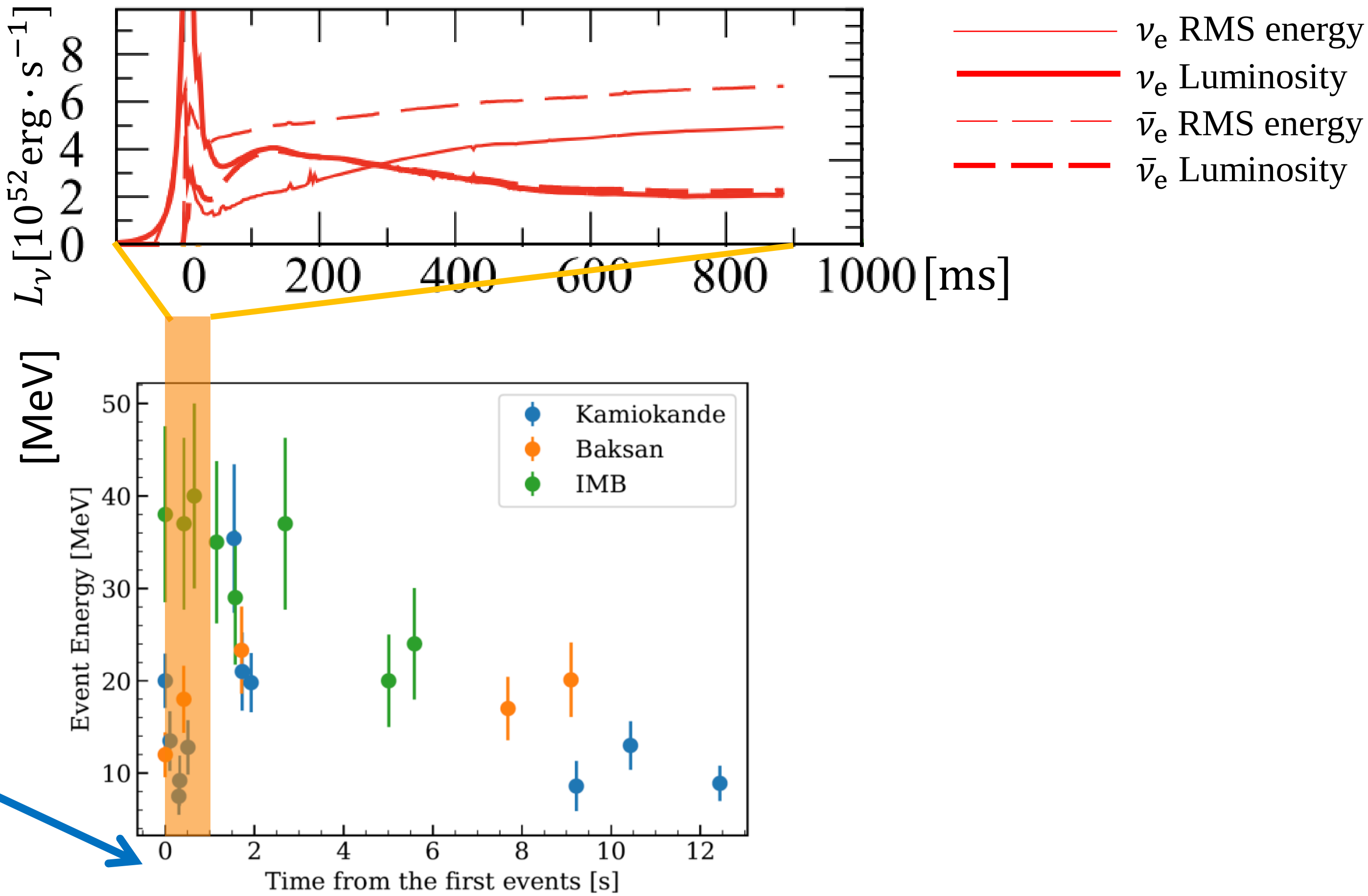
Supernova evolution



Supernova simulation problem

- Most simulations concentrate on early 1 sec.

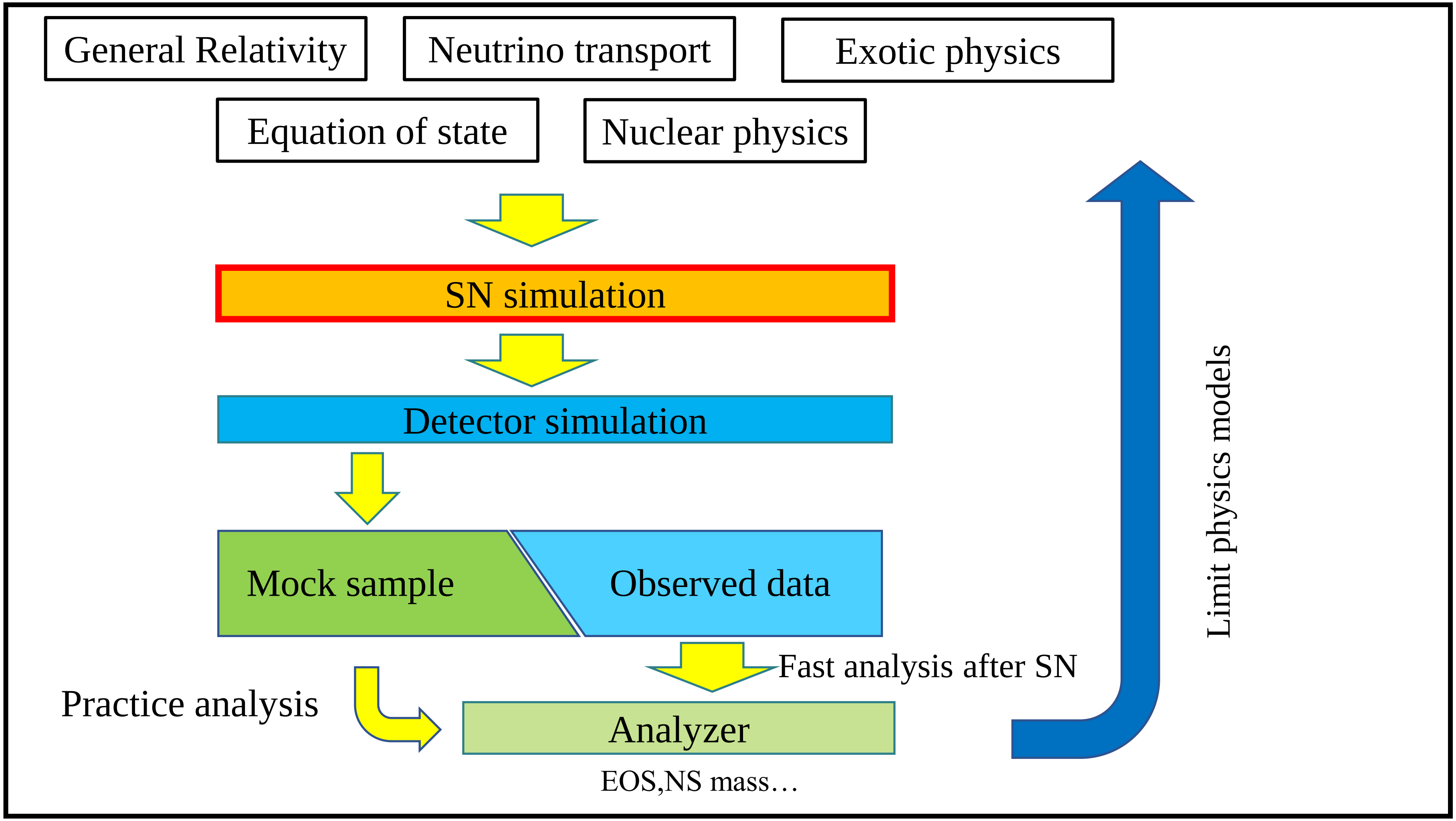
Example of simulation
Suwa et al. (2016)



We can compare theory and observation only for this time.

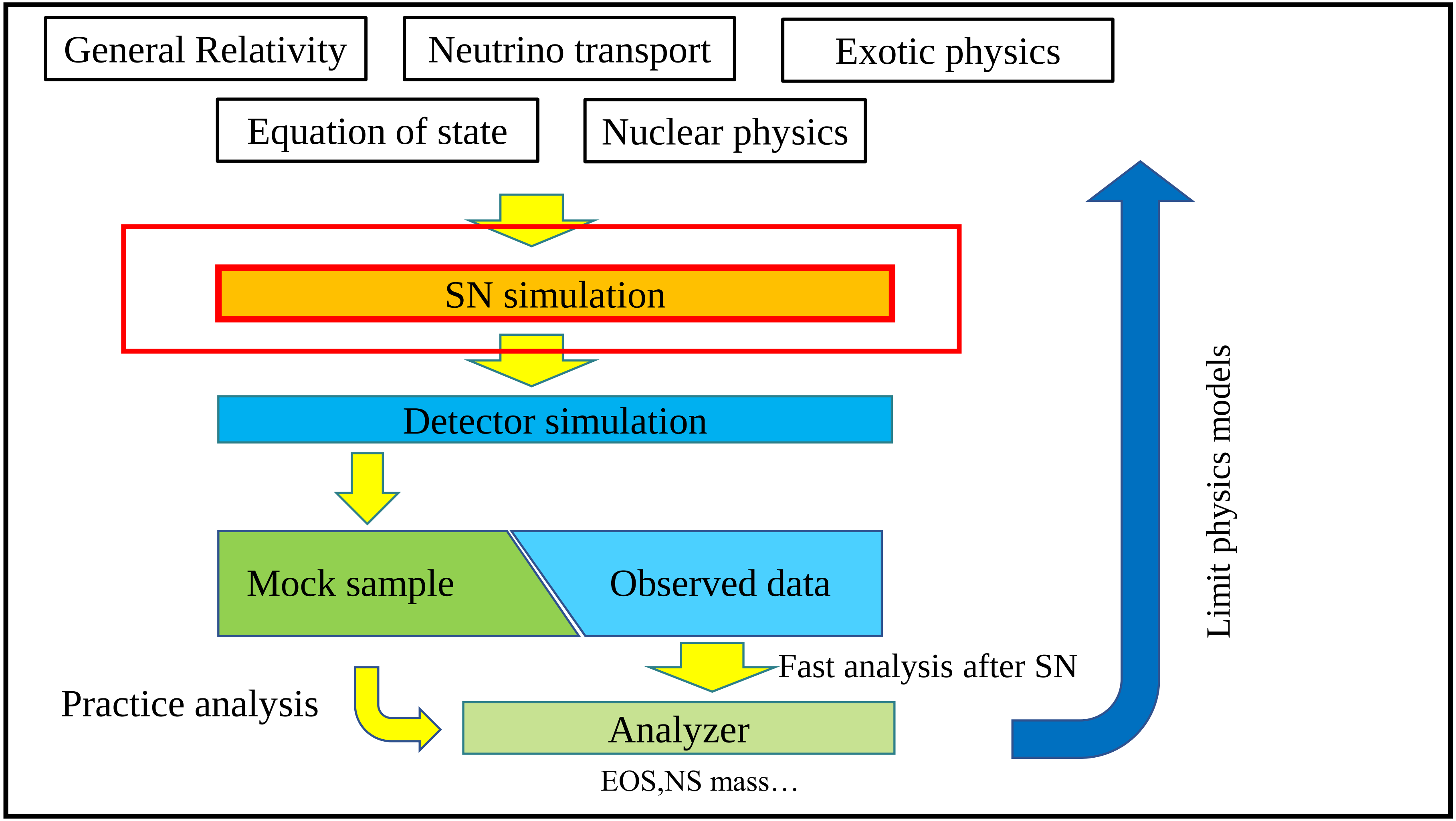
We developed a long-term simulation and are developing an analysis method.

Integrated framework



- Simulator which calculates from explosion to observation on earth.
- If a supernova is detected, the framework quickly analyze.

SN simulator



- Supernova simulation

Method of long time simulation

- Simulate supernovae in one-dimension
- Code
  - GR1Dv2 (public code: http://stellarcollapse.org)
    - ➢O'Connor, ApJS 219 24 2015
  - Gravity: General relativity
  - M1 scheme
  - Modified for long-term simulation
    - ➢Resolved reference out of physics tables
    - ➢Optimized resolution of time and space
  - EOS: DD2
  - Progenitor: $9.6M_{\odot}$, zero metallicity

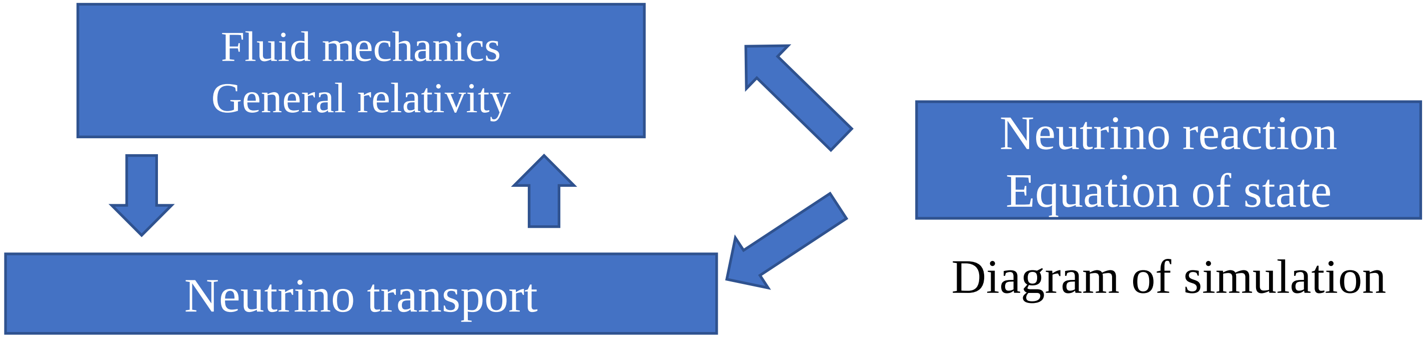
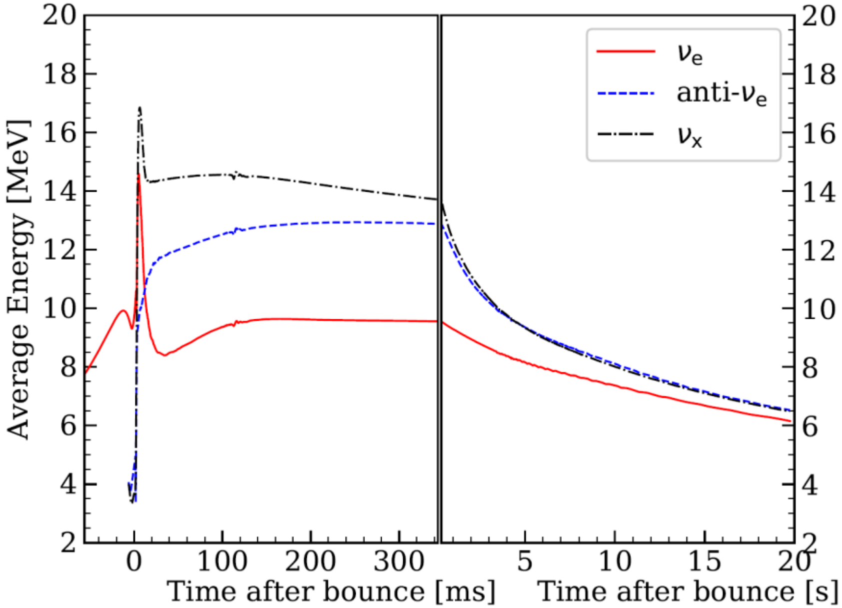


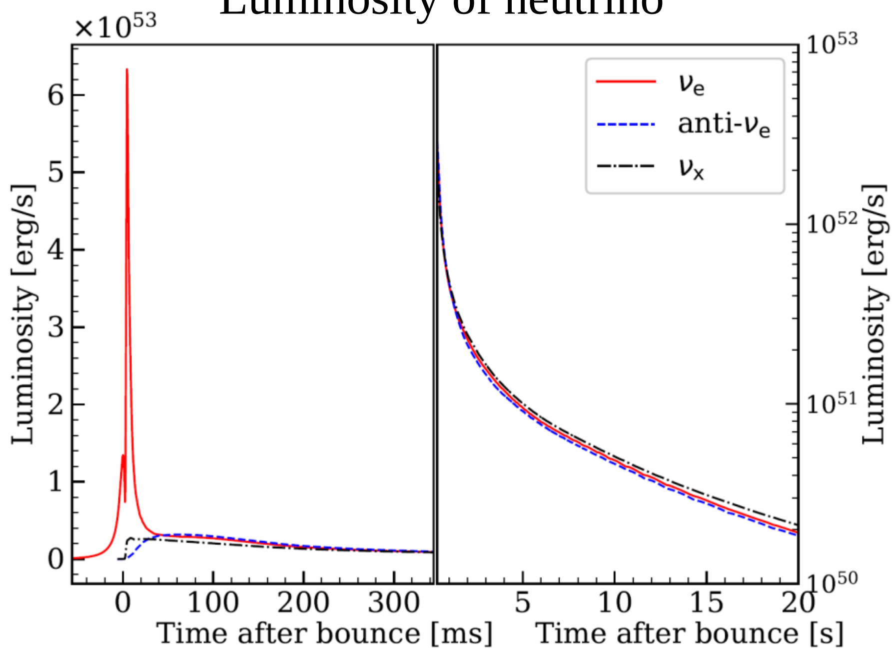
Diagram of simulation

Long-term neutrino emission

Average energy of neutrino



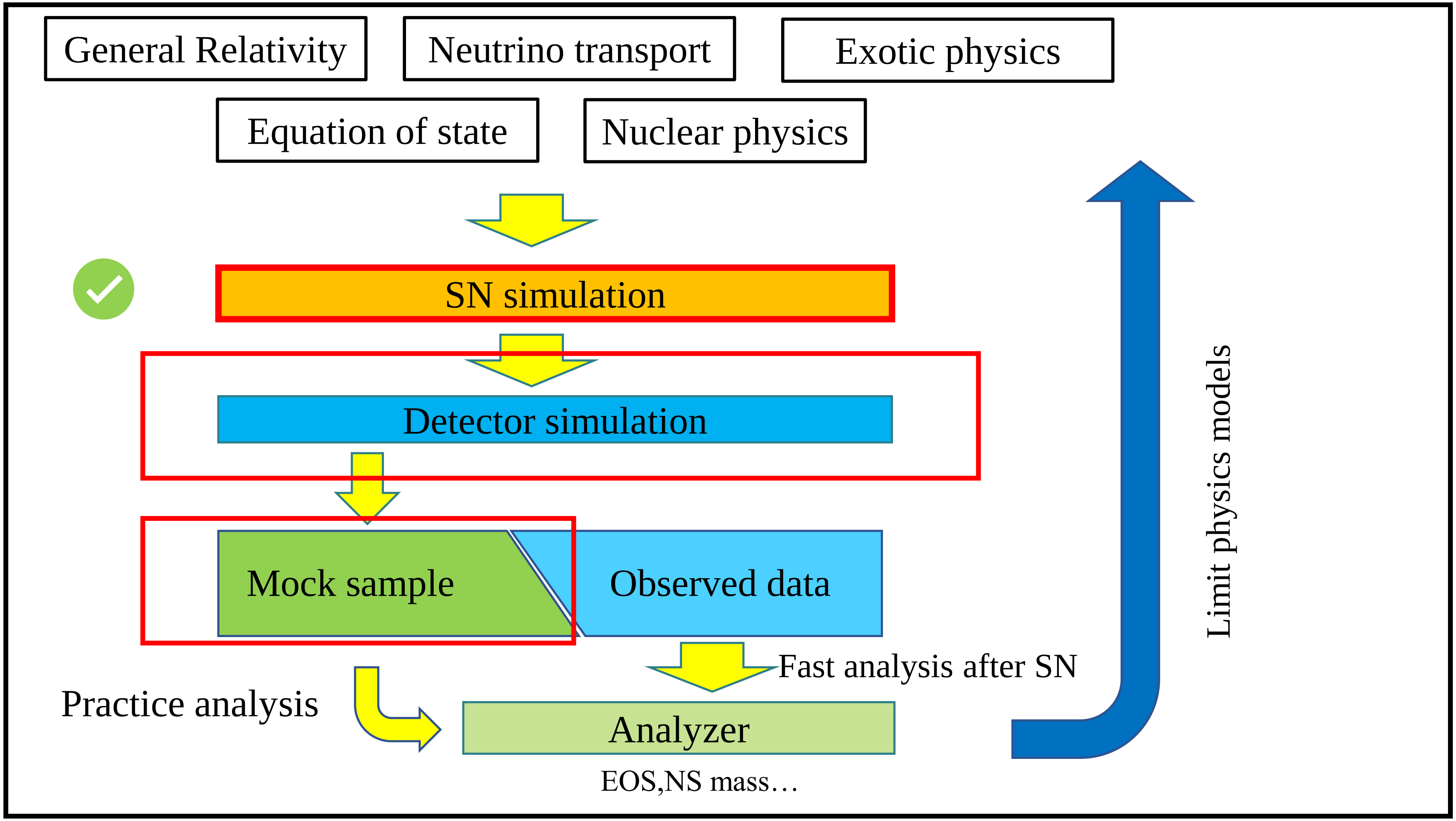
Luminosity of neutrino



ν_x : μ and τ (anti) neutrinos

- Average energies decrease from above 10 MeV to 6 MeV
- $\langle E_{\nu_e} \rangle < \langle E_{\bar{\nu}_e} \rangle < \langle E_{\nu_x} \rangle$
- Luminosities decrease from 10^{53} erg/s
  - ➢These features agree with other simulations.
  - ➢PNS cooling is calculated.

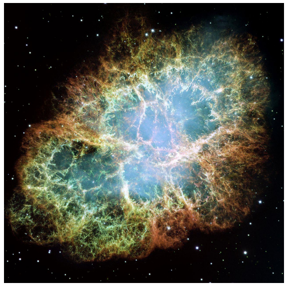
Detector simulator



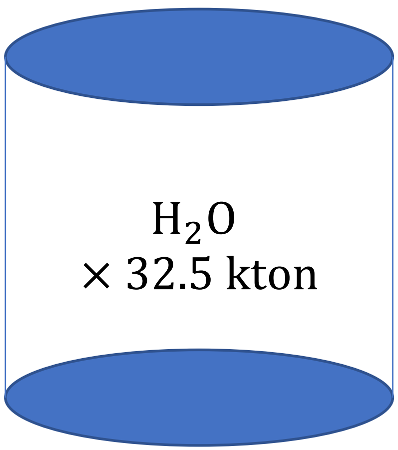
- Detector simulation
- Simulates how signals of supernovae look like on earth
- Mock Samples are used for analysis practice and detector evaluation.

Event simulation

Explosion



10 kpc



Super-Kamiokande(SK)

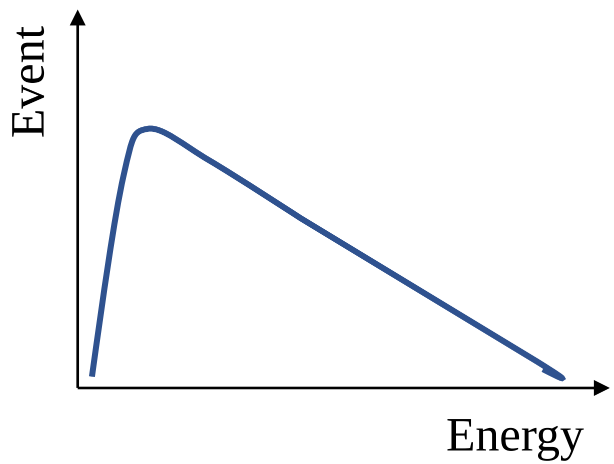
□ Reaction channel

Inverse Beta Decay(IBD)

- $\bar{\nu}_e + p \rightarrow e^+ + n$
- Amount: more than 90%
- Direction sensitivity ： No

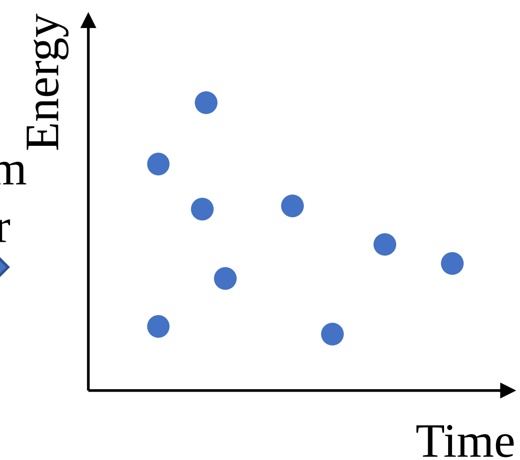
Electron scattering(ES)

- $\nu + e^- \rightarrow \nu + e^-$
- Amount: 1/20 of IBD
- Direction sensitivity ： Yes

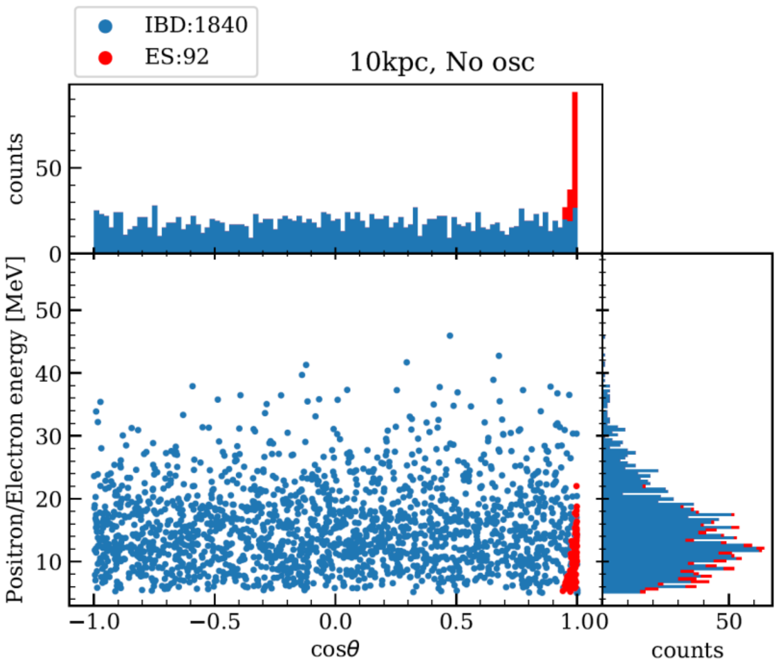


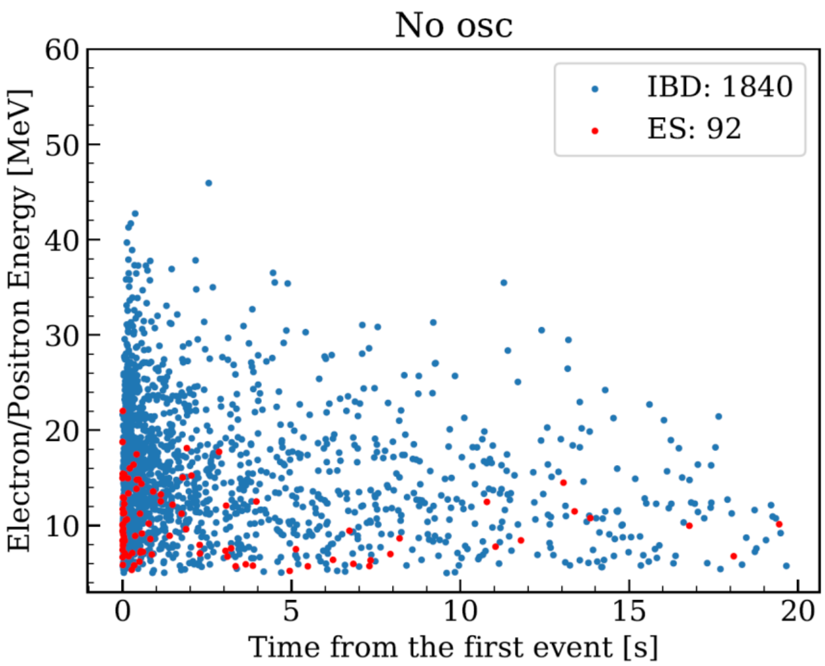
Event distribution per time

Random number



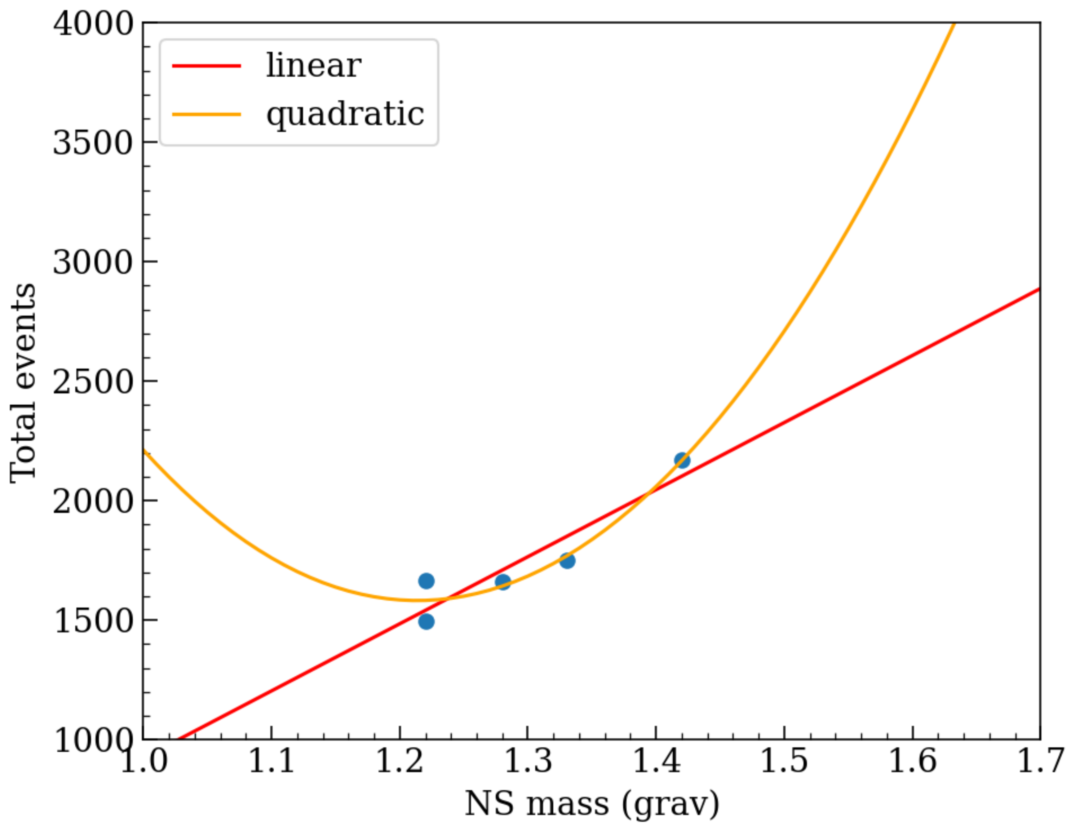
Scatter plot (Mock sample)





- Each event is simulated with random number (10kpc)
- Left：cosine distribution between neutrinos and charged leptons.
- Right：Time evolution of energy
  - Almost all IBD, ES scatters forward.

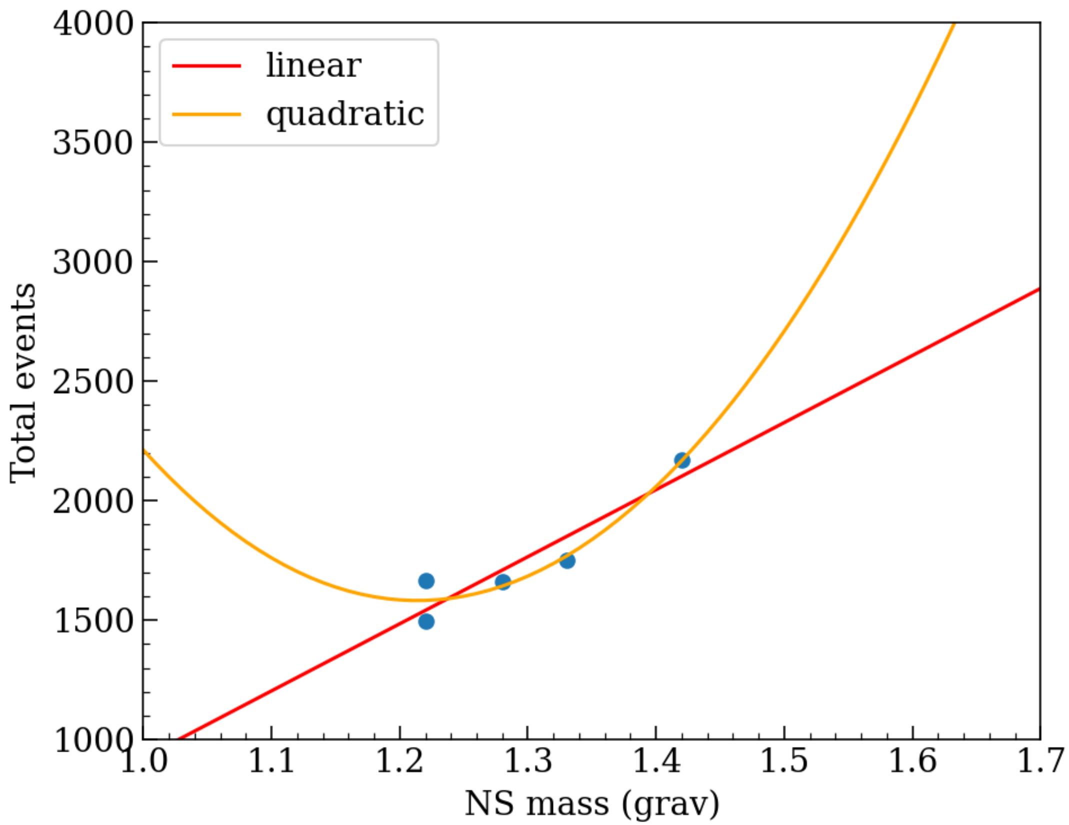
Total events and neutron star masses



Simply put,
total events may be proportion to $\frac{GM^2}{R}$?

- 5 simulations which reached to 10 seconds.
- Vertical axis: the number of events at SK for 10 s.
- Horizontal axis: neutron star mass
- It seems that heavier neutron stars lead to more neutrino events.
  - It possible to estimate neutron star mass from neutrino events.

Total events and neutron star masses



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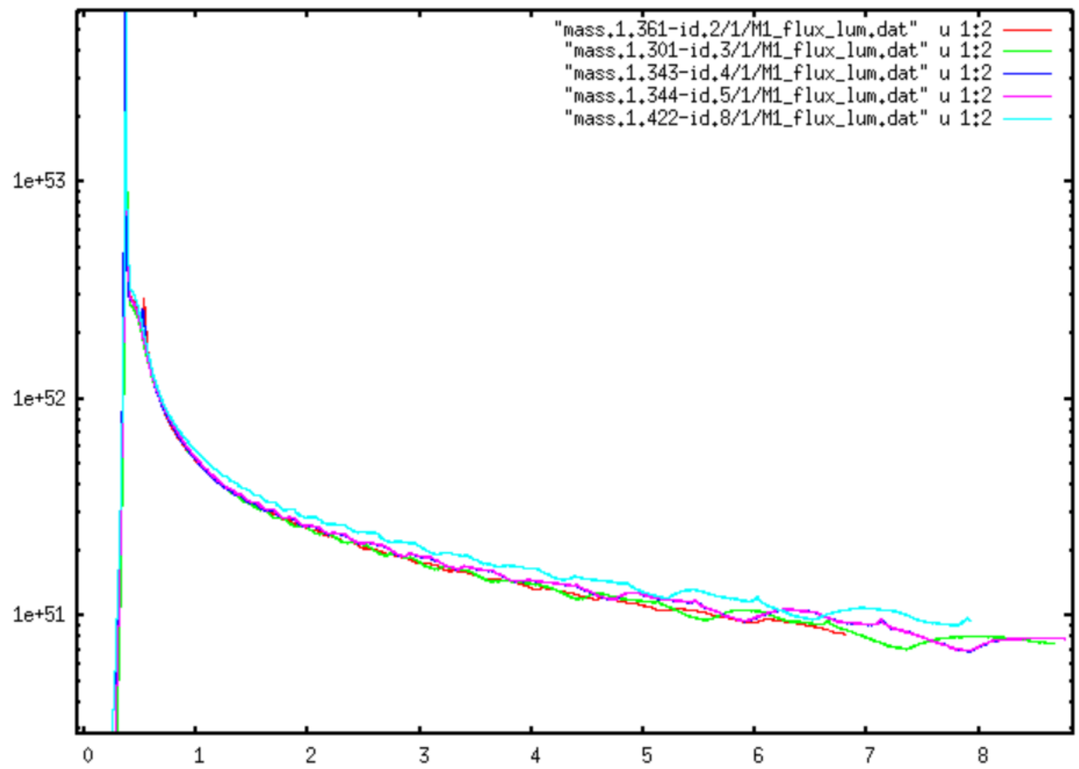
Calculating many progenitors



- We have 50 progenitors, which explode in 1D.

Gravitational eigenmodes

- Supernovae also emit gravitational waves.
- We estimated eigenmode frequencies with the asteroseismology approach.
  - Used GREAT (https://www.uv.es/cerdupa/codes/GREAT/)
  - Torres-Forné et al, MNRAS, 474, 5272 (2018)
  - Torres-Forne et al, arxiv:1806.11366 (2018)
- Linear perturbation analysis both of fluid and metric.
- Calculated eigenmodes from the result of the $9.6M_{\odot}$ progenitor in post-process

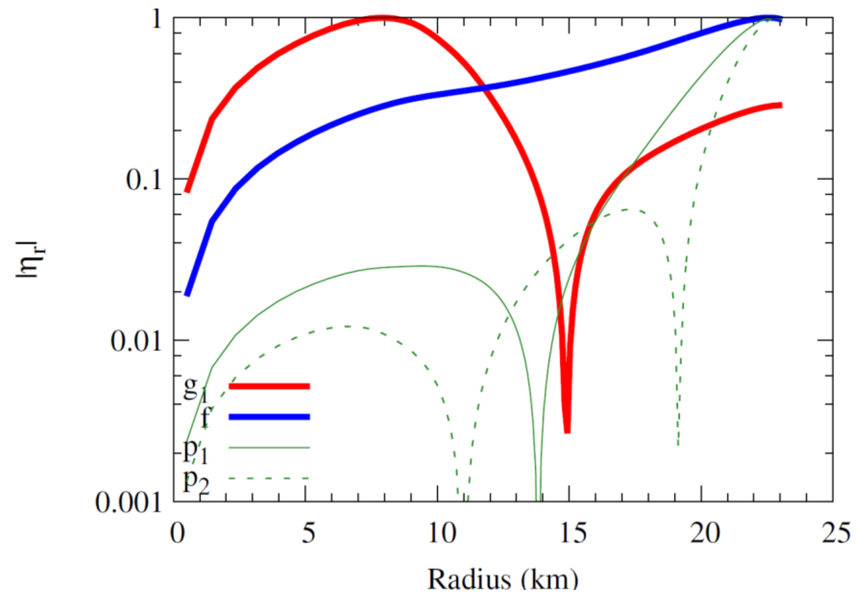
- Linear perturbation equations

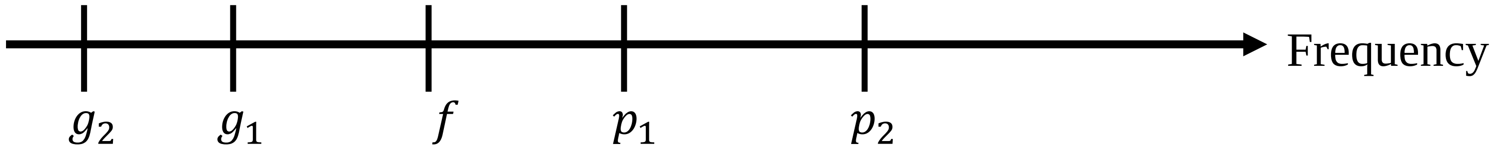
$$\partial_r \eta_r + \left[\frac{2}{r} + \frac{1}{\Gamma_1}\frac{\partial_r P}{P} + \frac{\partial_r \psi}{\psi}\right]\eta_r + \frac{\psi^4}{\alpha^2 c_s^2}\left(\sigma^2 - \mathcal{L}^2\right)\eta_\perp = \frac{1}{c_s^2}\frac{\delta\hat{Q}}{Q} - \left(6 + \frac{1}{c_s^2}\right)\frac{\delta\hat{\psi}}{\psi},$$

$$\partial_r \eta_\perp - \left(1 - \frac{\mathcal{N}^2}{\sigma^2}\right)\eta_r + \left[\partial_r \mathrm{ln} q - \mathcal{G}\left(1 + \frac{1}{c_s^2}\right)\right]\eta_\perp = \frac{\alpha^2}{\psi^4\sigma^2}\left[\partial_r(\mathrm{ln}\rho \mathrm{h})\left(1 + \frac{1}{c_s^2}\mathcal{G}\right)\right]\left(\frac{\delta\hat{Q}}{Q} - \frac{\delta\hat{\psi}}{\psi}\right),$$

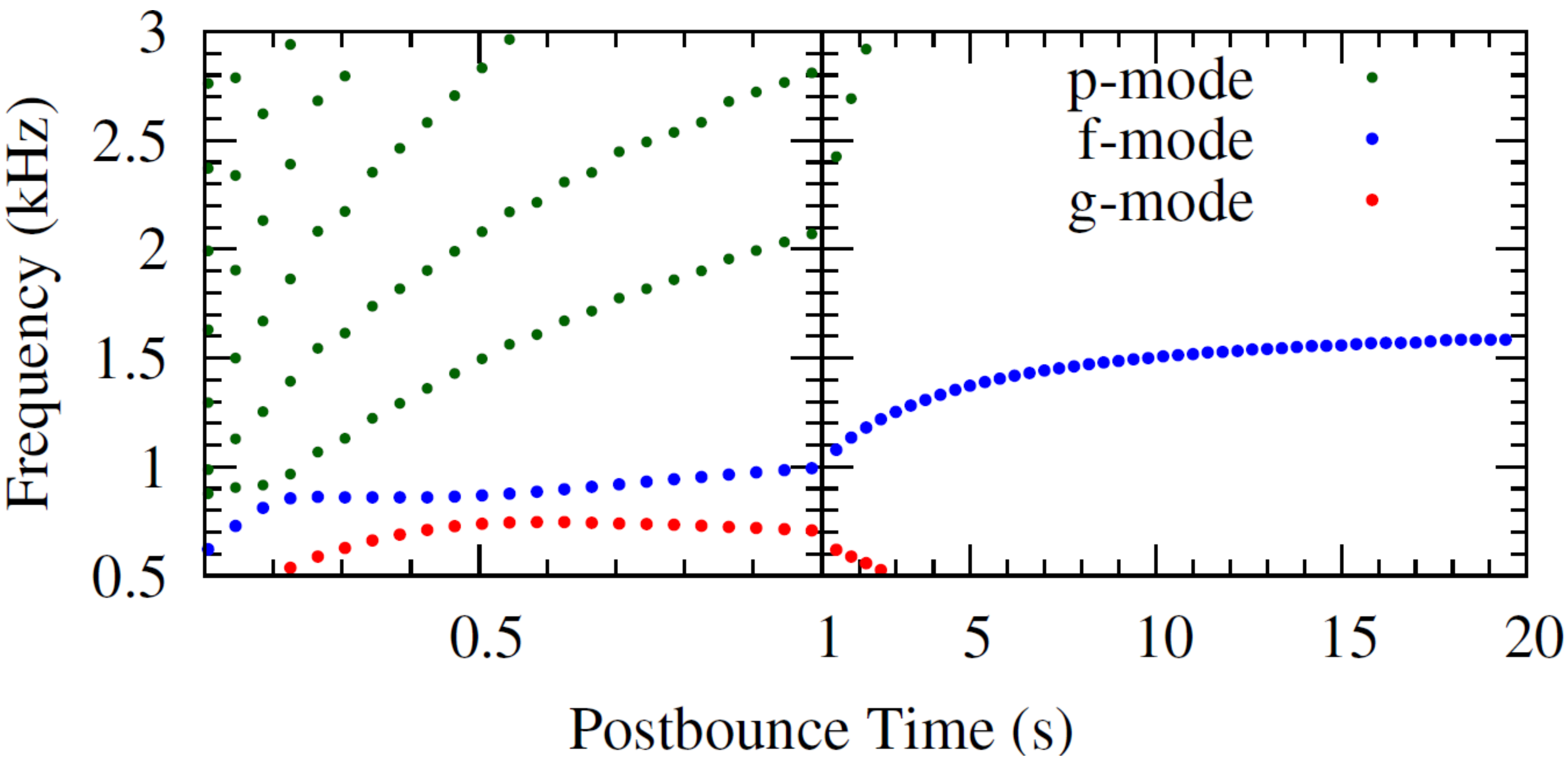
Kinds of eigenmode

- p_i-mode
  - "i" is the number of nodes
  - Restoring force: pressure
  - Frequencies increase as nodes increase
- f-mode
  - Fundamental mode of the p-mode
- g_i-mode
  - Restoring force: buoyancy
  - Frequencies decrease as nodes increase



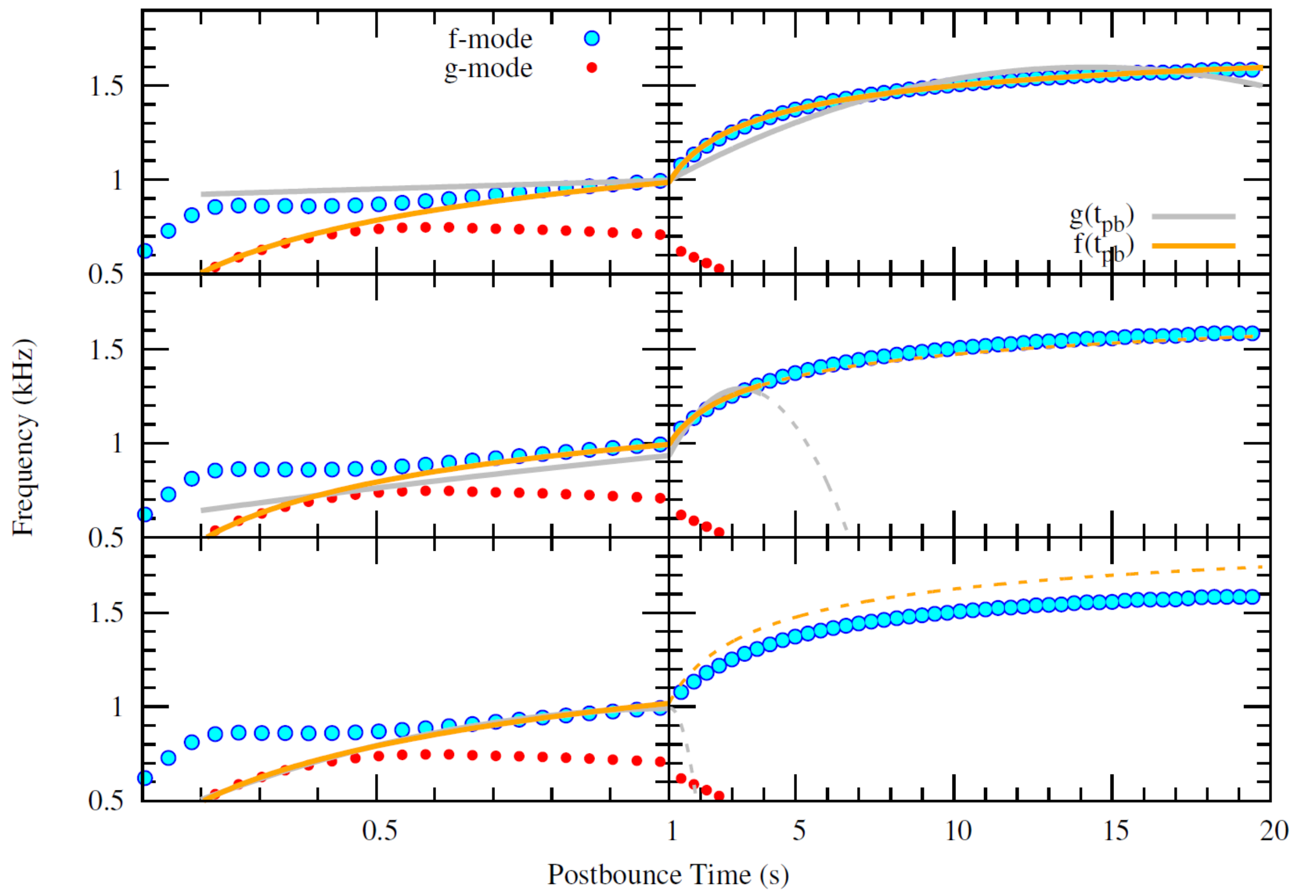


Gravitational wave frequency



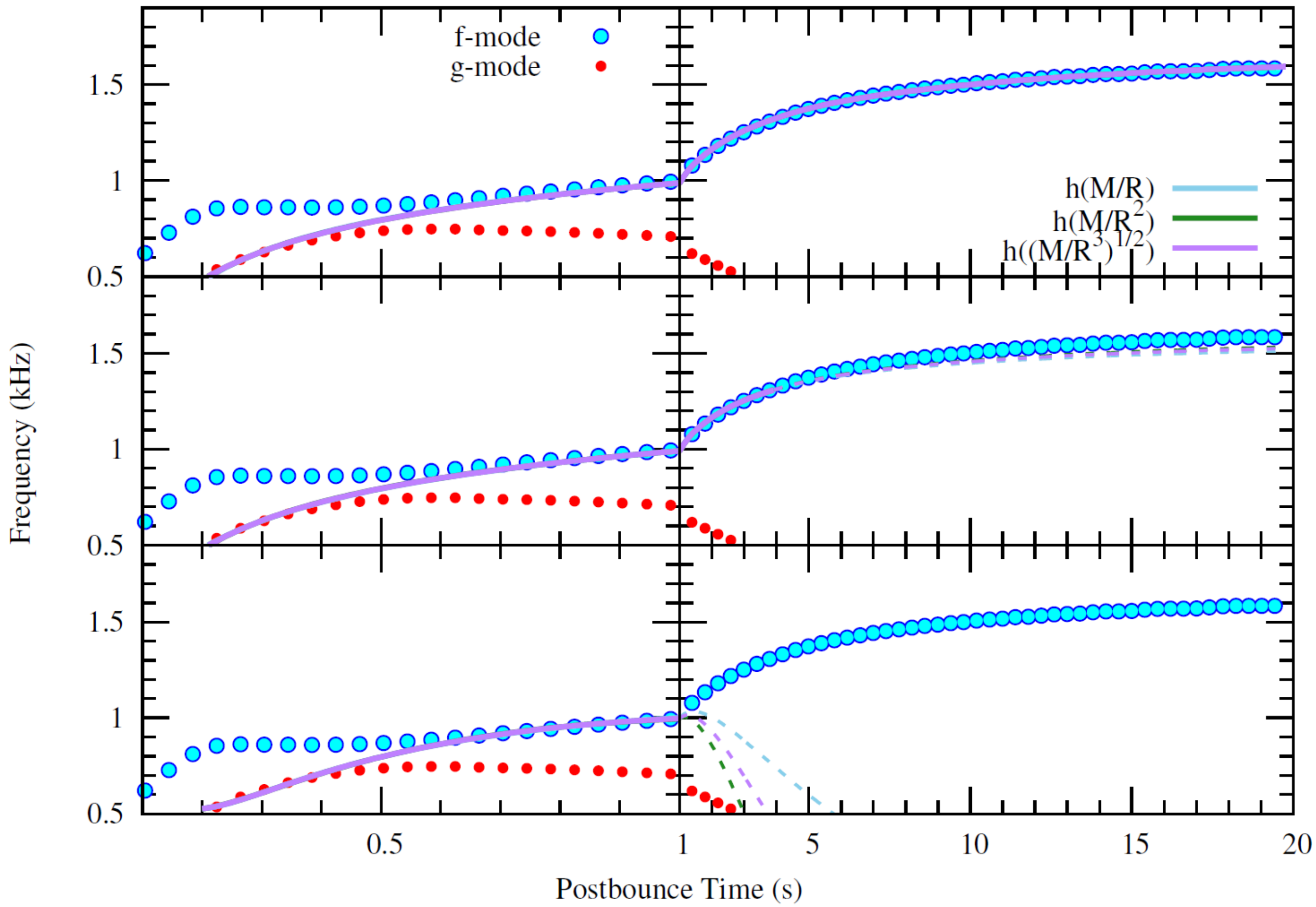
- Calculated eigenfrequencies up to 20 seconds.
- Differences of frequencies increase with time.
- Next, we make fitting functions.

Gravitational wave fitting



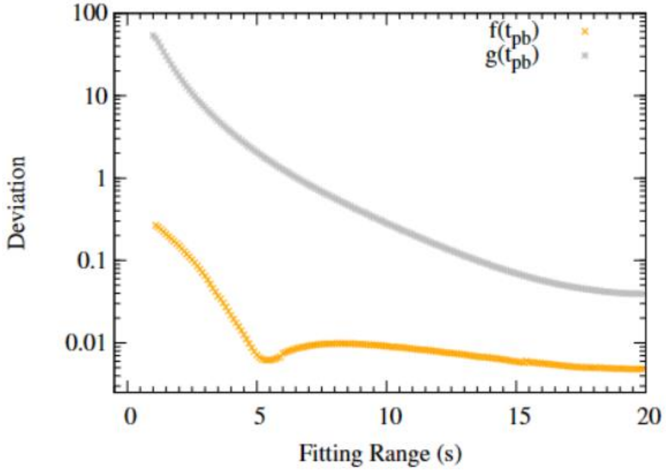
- Fitted with postbounce time
- g(x): quadratic function
- $f(x = t_{\mathrm{pb}}) = \dfrac{a_1 x^{a_4}}{x^{a_4} + a_2} + a_3,$

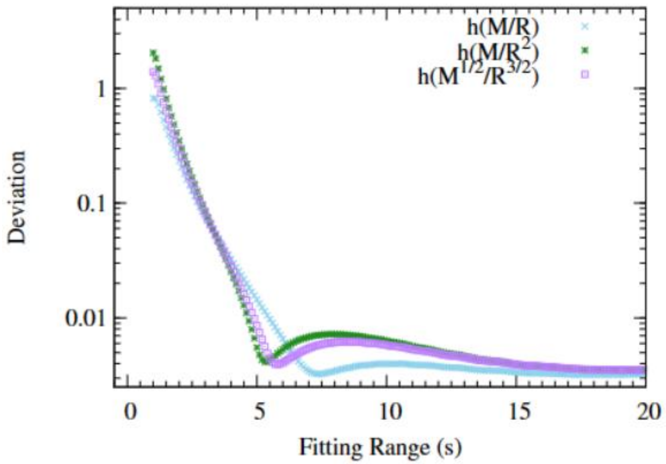
Gravitational wave fitting with mass and radius



- Fitted with mass and radius
- $h(x) = c_1 + c_2 \log(x) + c_3 x + c_4 x^2,$
- x: $M/R,\ M/R^2$ and $\sqrt{M/R^3}$

Fitting accuracy





- Deviation
  - $$D(t_{\rm end}) \equiv \int_{T_{\rm start}}^{T_{\rm sim}} \left| \frac{A_{\rm sim}(t) - A_{\rm fit}(x(t),\ t_{\rm end})}{A_{\rm sim}(t)} \right| dt \, / \, (T_{\rm sim} - T_{\rm start}) \,,$$

Summary

Summary

- Supernovae emit neutrinos and gravitational waves.
- Established the long-term neutrino radiation hydro simulation
- Estimated neutrino signals at Super-Kamiokande
  - We may be able to estimate neutron star mass via neutrino events.
- Estimated long-term gravitational wave eigemodes

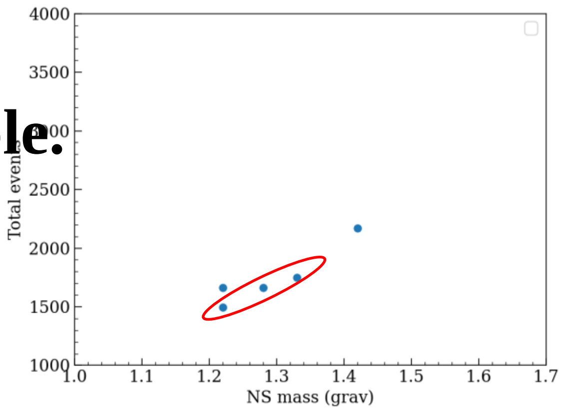
To do

- More progenitors will be simulated
- Develop an analysis method

My three simulations circled in red are available.

They reach to 20 seconds.

https://zenodo.org/record/5825648



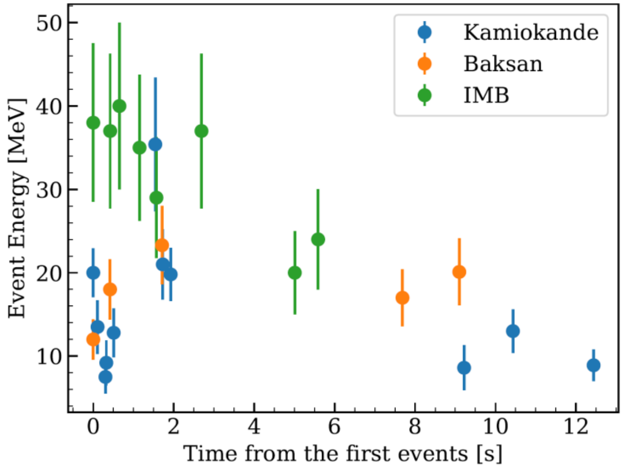
Back up

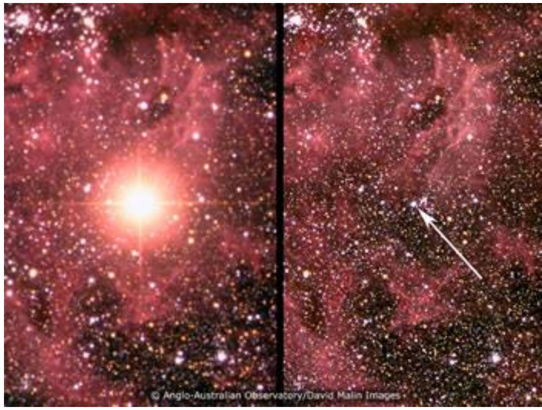
Supernova

- 8 times heavier stars than the sun happen huge explosion
- Complicated phenomenon in which all the four forces of nature are related
  - ➢Not analytic calculation but heavy computation is needed
- Energy of 10^{53}erg is released as neutrinos.
  - ➢ Only one observation in 1987 (SN1987A)

SN1987A information

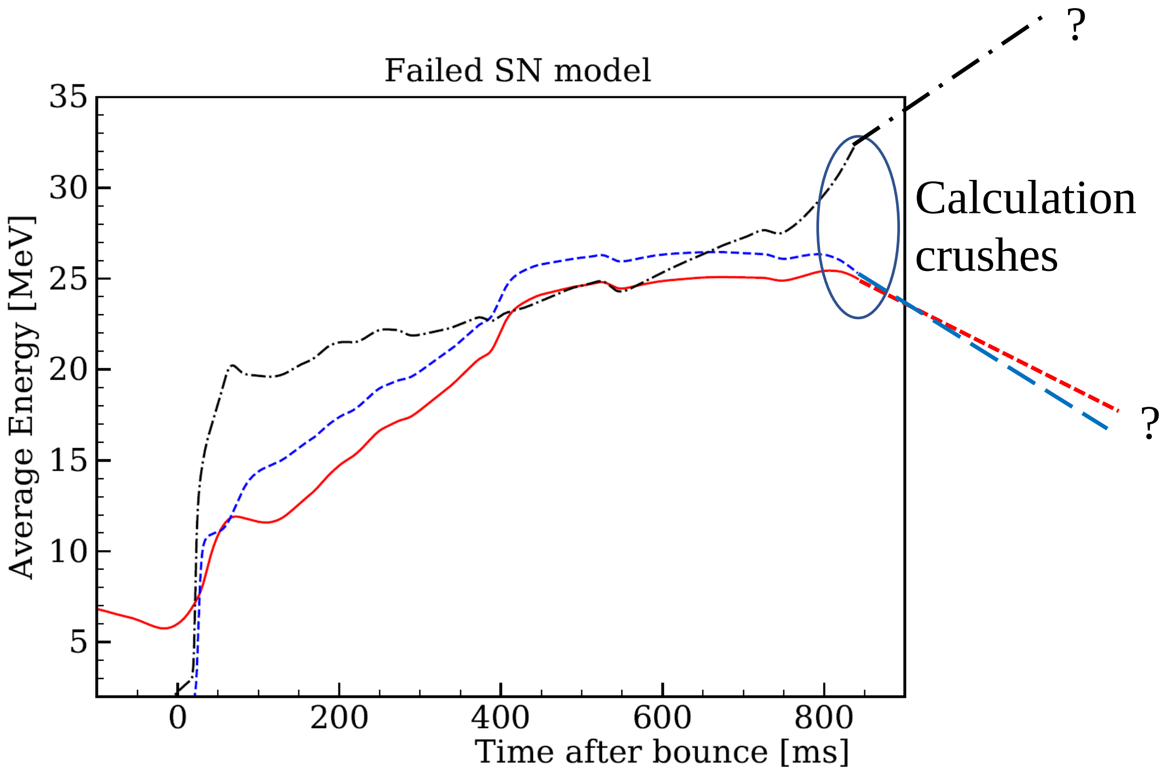
Distance:51.2 kpc
Number of events: Detector
- 11: Kamiokande(2.14 kton)
- 8: IMB [2]
- 5: Baksan [3]





[1]Hirata et al. 1987
[2]Bionta et al. 1987
[3]Alekseev et al. 1987

First idea



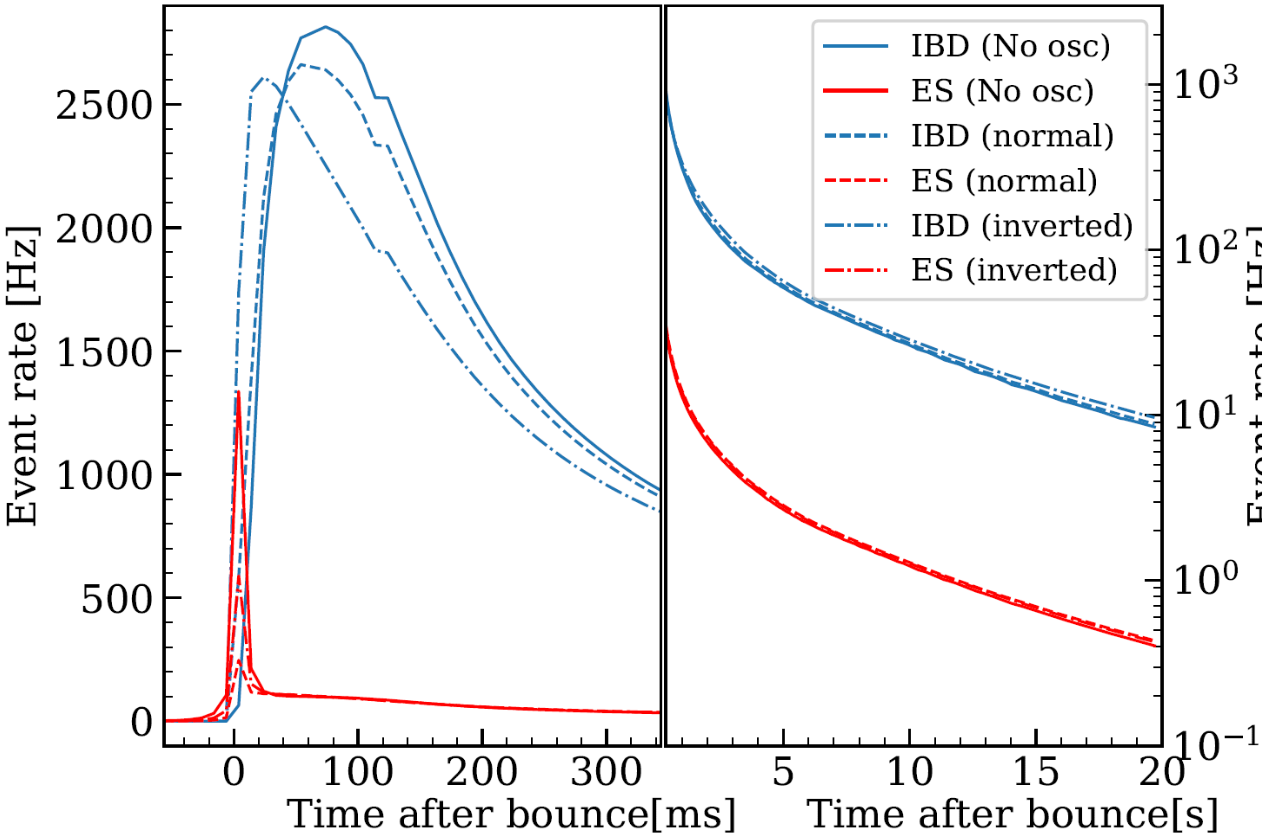
Misner-Sharp metric

$$ds^2 = -e^{2\phi}dt^2 + X^2dr^2 + d\Omega^2$$

$$X = 1/\sqrt{1 - \frac{2M}{R}}$$

- Calculation in case of black hole formation is more difficult
- Because metric diverges at an event horizon.

Reaction rate



- Assumed a supernova happen at 10 kpc（Distance to the galactic enter: 8kpc）
- About 2,000 events at 20 seconds
- In the later time, neutrino oscillation has little influence

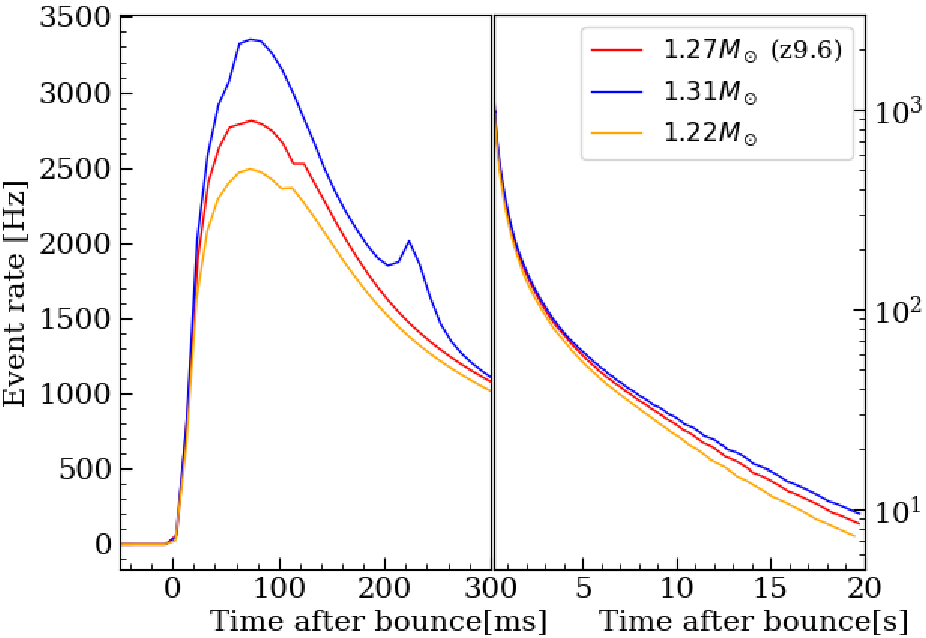
Recent simulations

	Huedepohl (1D)	Fischer (1D)	Multi-dimension Takiwaki(2016), Suwa(2016)… etc	This study
Iron core	×	○	○	**○**
Natural explosion	○	×	○	**○**
Max time	20 s	20 s	< 1 s	**20 s**

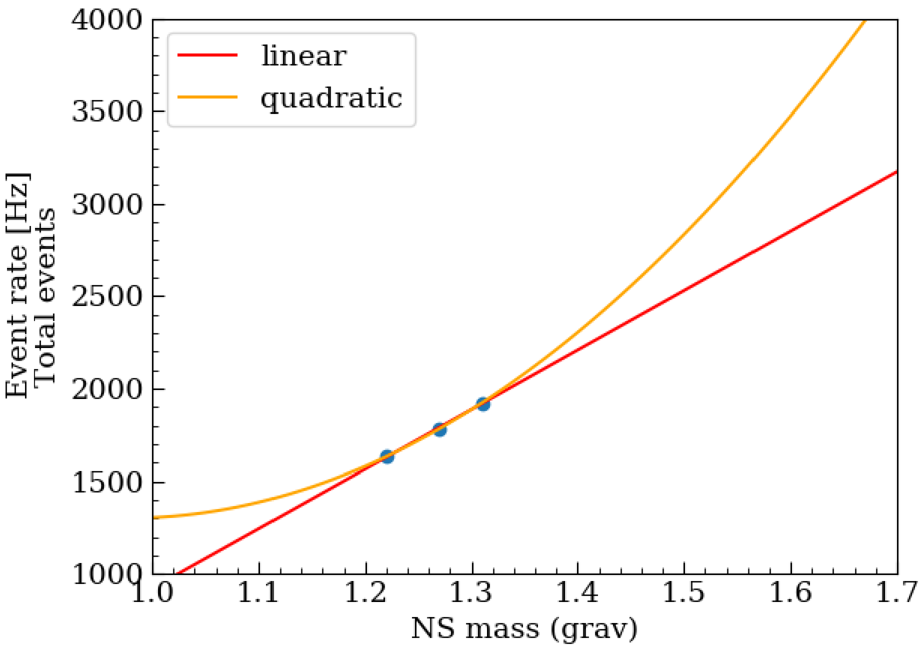
- To explode without artificial methods in one-dimension is difficult
  - ➢ Enhancement of neutrino reaction rates
  - ➢ Removal of material accreting
- Long time simulation in multi-dimension is impossible
- We do long time simulation in one-dimension without artificial methods

Neutrino and neutron star mass

Event rate at 10 kpc



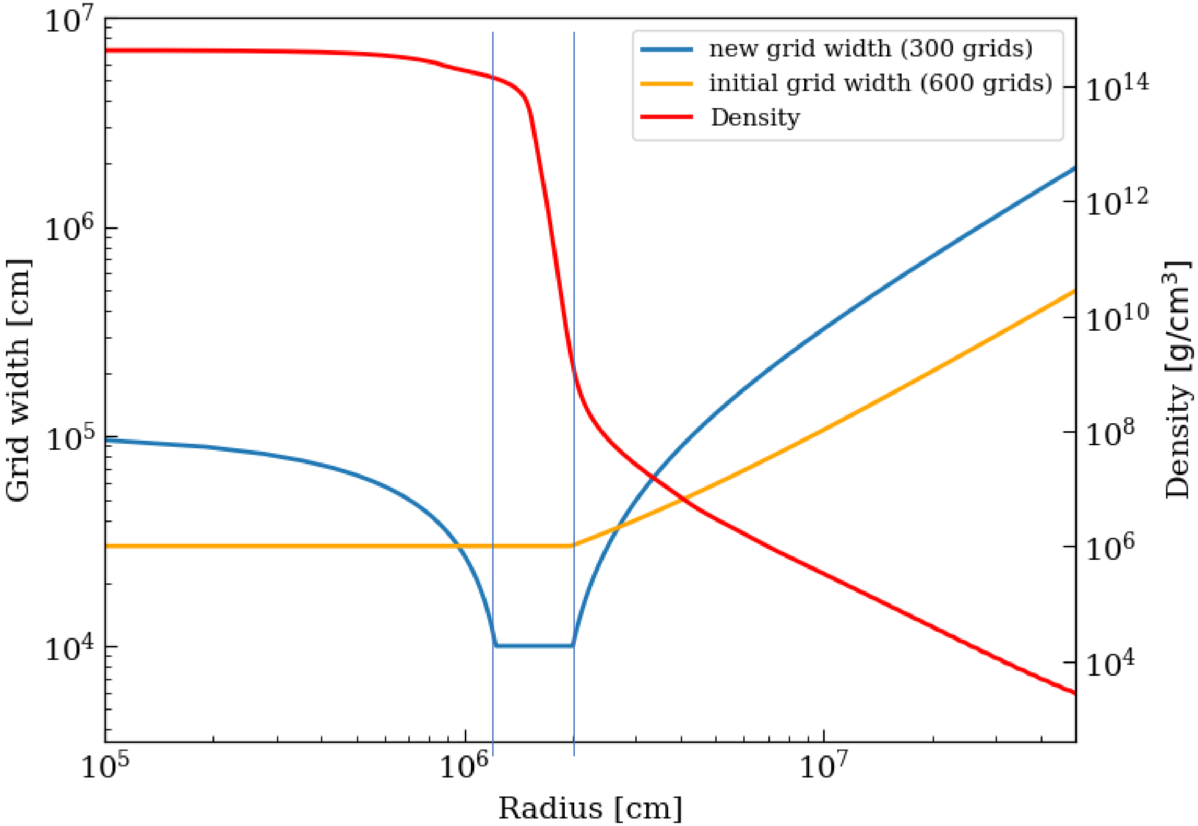
Relation between the number of events and neutron star mass

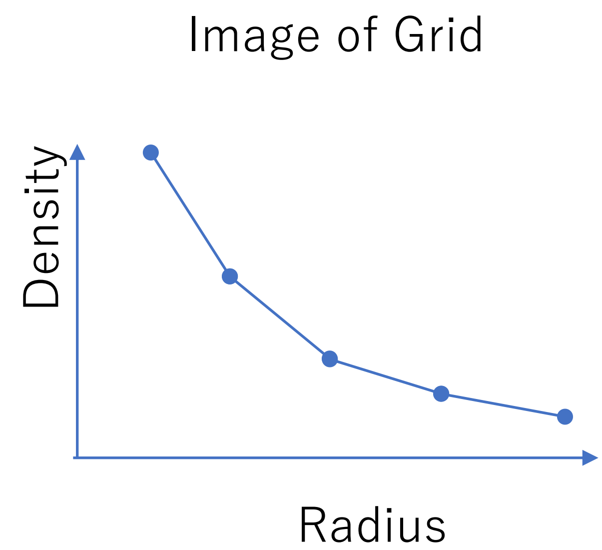


- Three simulations which lead to different neutron star mass
- If distance is determined, neutron star mass is maybe determined.
  - ➢More simulations are needed.

➢In addition, I’m developing simulation in the case of BH formation.

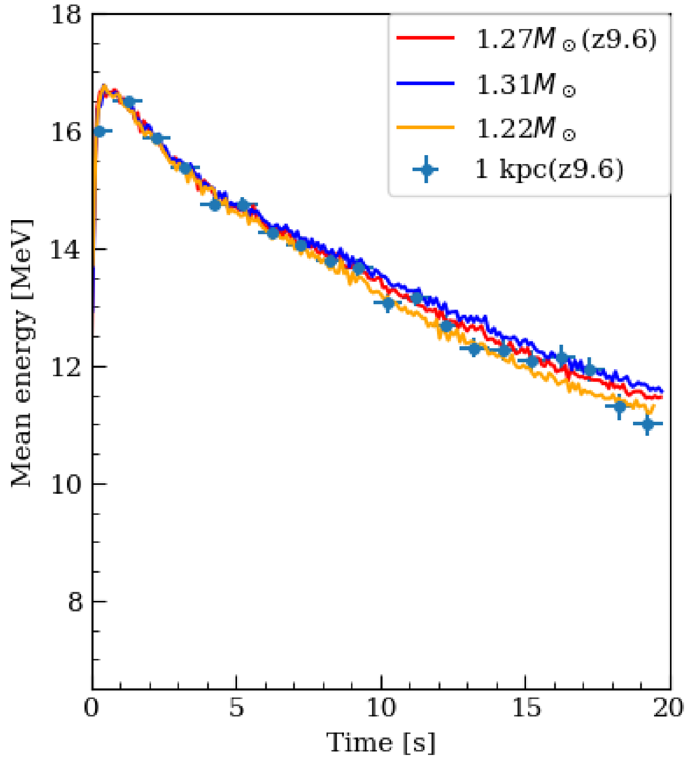
Device of grids

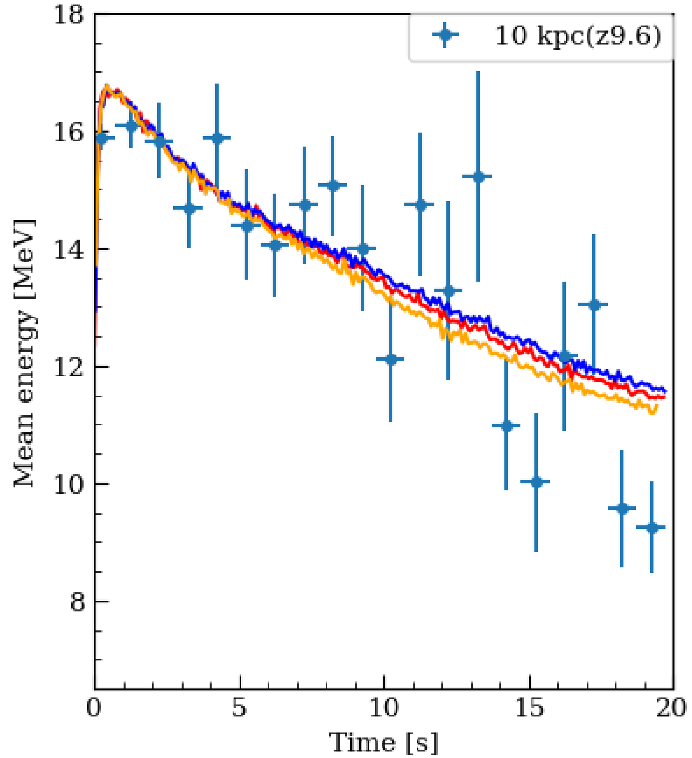




- Red：Density structure of PNS
- Yellow：Initial grids (600 grids)
- Blue：Optimized grids (300 grids)
- The region in which the density drastically changes is finely resolved.
  - ➢Initial grids make calculation stop at about 5 sec.
  - ➢Cost is also too high

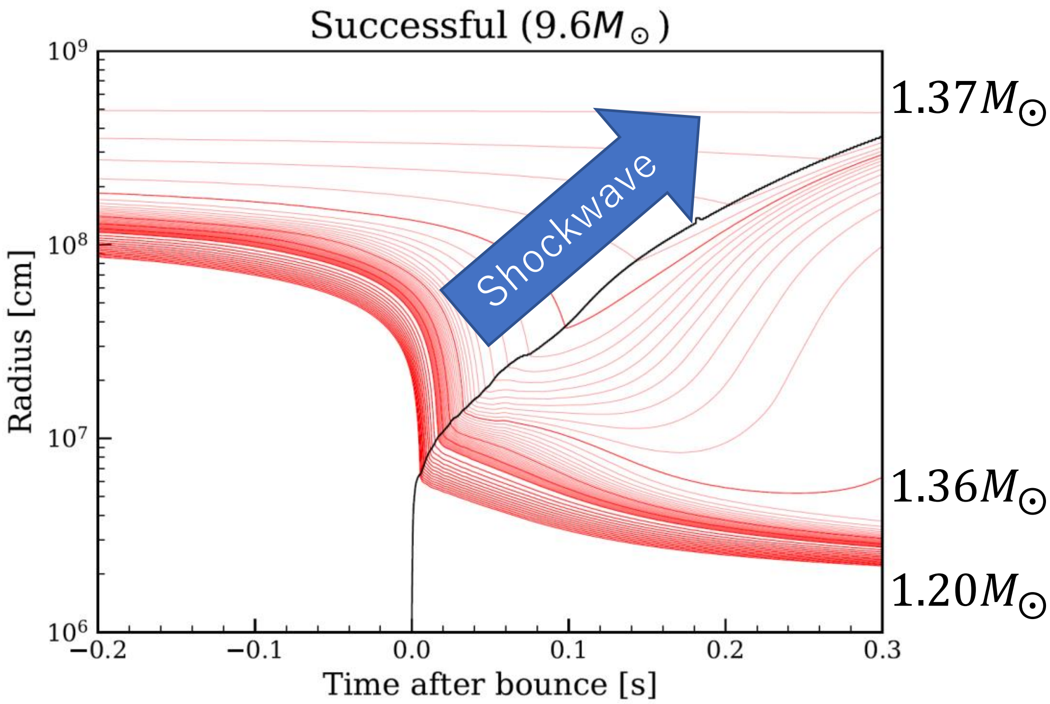
平均エネルギーの発展

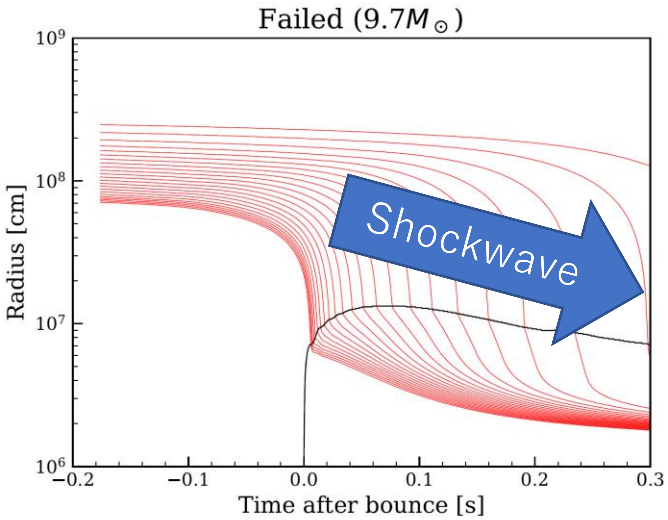




- **実線**:無限個のイベントの平均エネルギー
- **マーカー**:有限個のイベントの平均エネルギー(z9.6)
  - ➢エラーバー: $\sqrt{\dfrac{1/N_{\mathrm{bin}} \times \sum_{i=1}^{N_{\mathrm{bin}}} (E_i - \bar{E})^2}{N_{\mathrm{bin}}}}$
- 個数だけでなく、エネルギー情報も使った比較が可能
- エネルギーの時間発展からのモデルの分別を目指す。

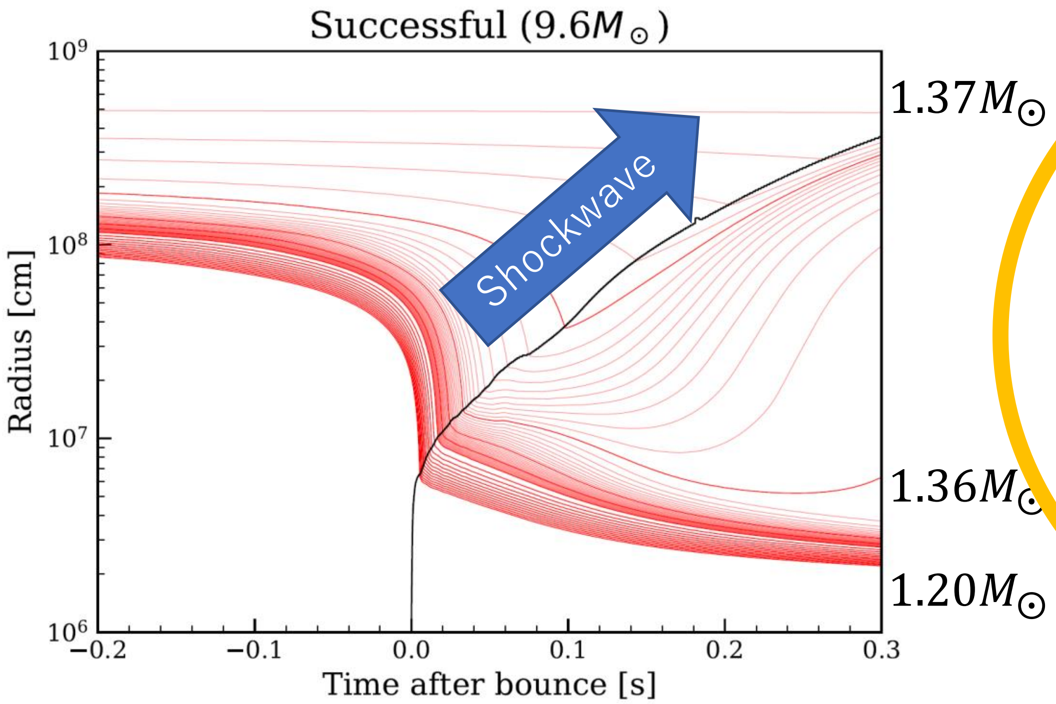
Light progenitor

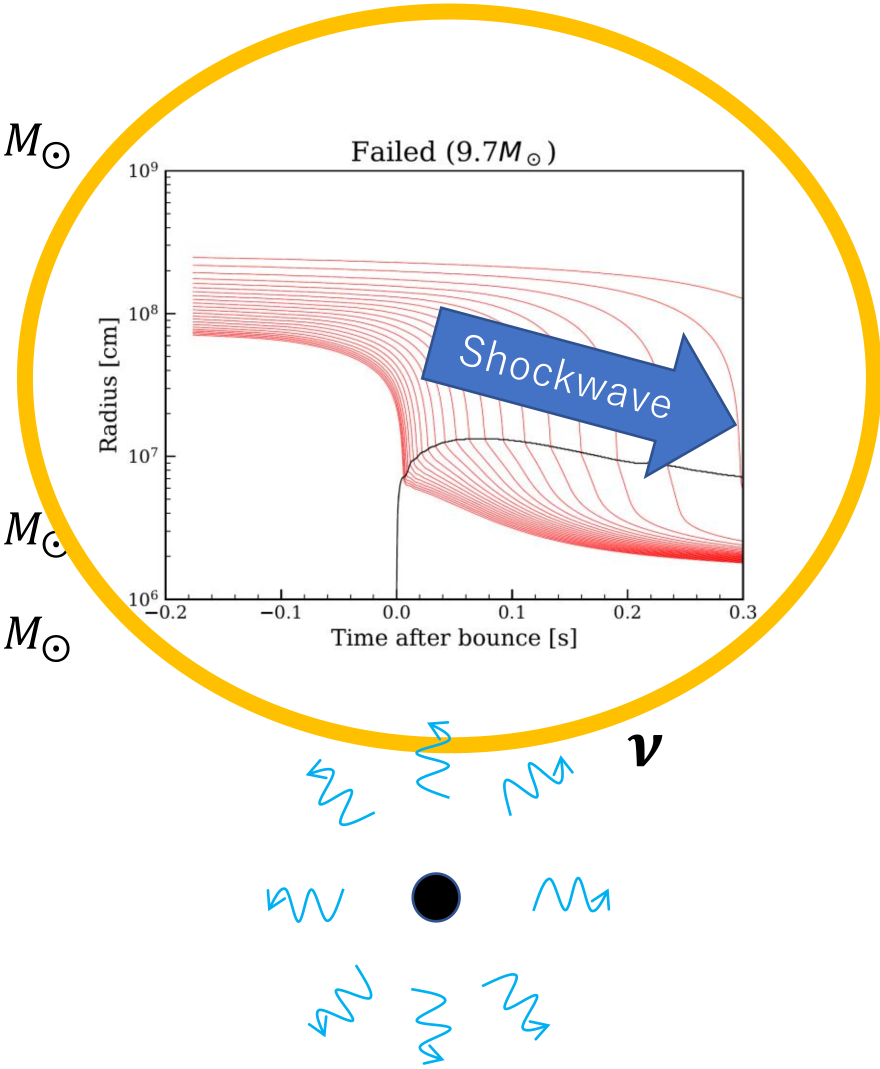




- Red：Radii at which densities are constant
- Black：Radius of a shockwave
- Succeed to explode with the suitable choice of progenitors and **without artificial methods**
  - ➢9.6 solar mss, initial metallicity is 0
  - ➢Called z9.6

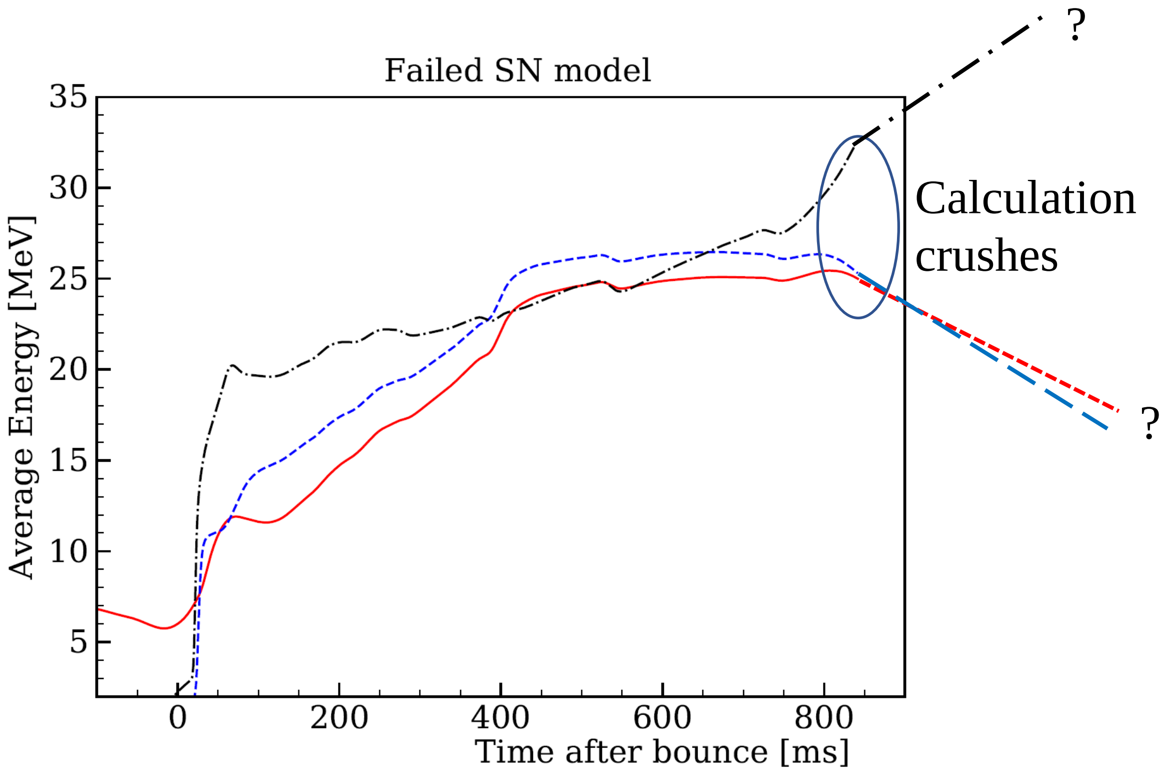
Black hole formation





- I want to also calculate the case of failed supernovae and black hole formation.

Calculation crush



Misner-Sharp metric

$$ds^2 = -e^{2\phi}dt^2 + X^2dr^2 + d\Omega^2$$

$$X = 1/\sqrt{1 - \frac{2M}{R}}$$

- Calculation in case of black hole formation is more difficult
- Because metric diverges at an event horizon.

Hernandez–Misner metric

- Misner-Sharp metric

$$ds^2 = -e^{2\phi}dt^2 + X^2dr^2 + d\Omega^2$$

- Introduce new time u

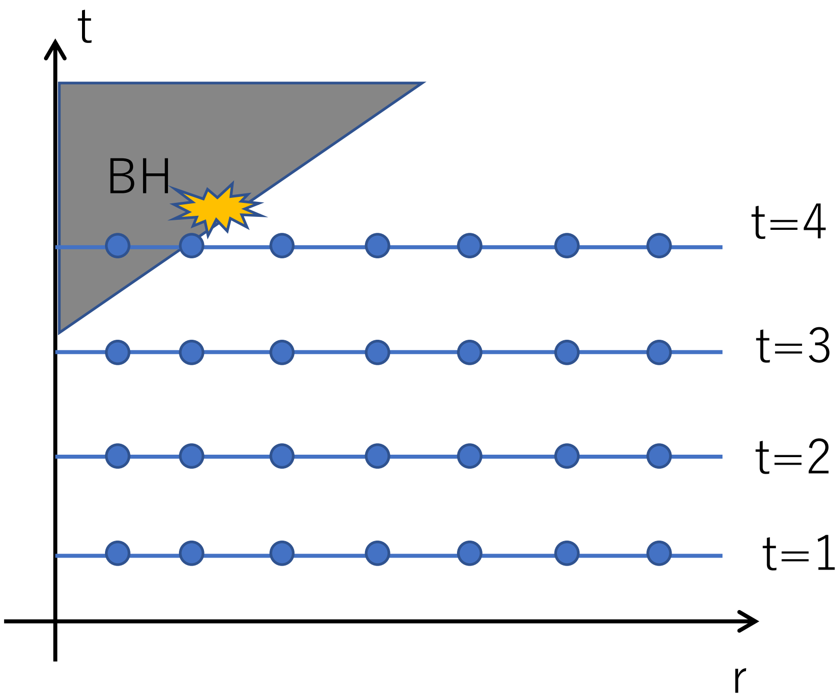
$$e^{\psi}du = e^{\phi}dt - Xdr$$

- Hernandez-Misner metric

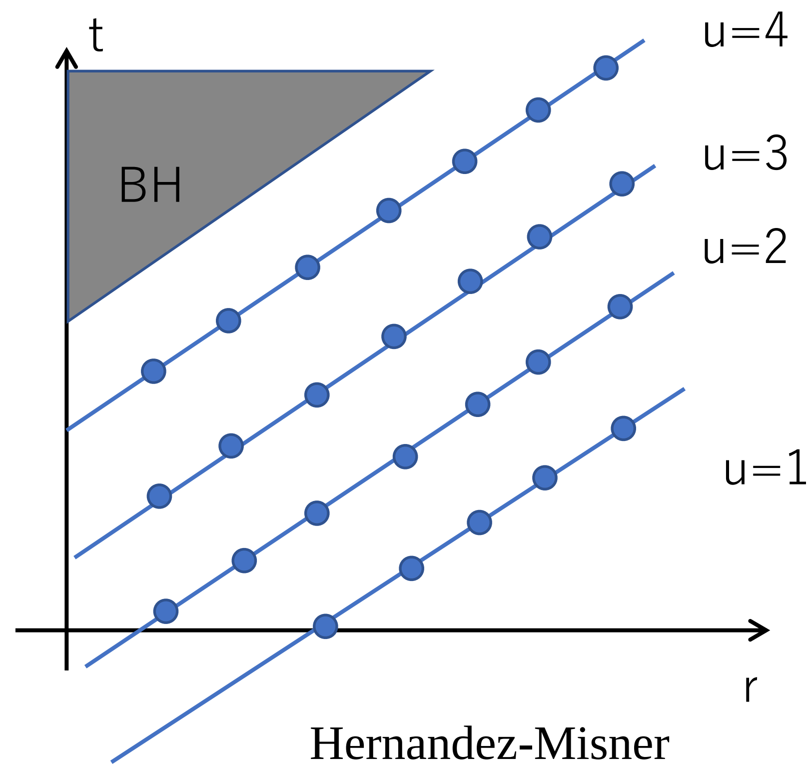
$$ds^2 = -e^{2\psi}du^2 - 2e^{\psi}Xdrdu + d\Omega^2$$

- The “u” is called observer time.

Difference between two metrics



Misner-Sharp



Hernandez-Misner

- Evolute time so that it avoids a black hole surface.
  - ➢Time is slower, closer to the center.